



THE UNIVERSITY *of* EDINBURGH

Edinburgh Research Explorer

Giant radio galaxies in the LOFAR Two-metre Sky Survey-I

Citation for published version:

Dabhade, P, Rottgering, HJA, Bagchi, J, Shimwell, TW, Hardcastle, MJ, Sankhyayan, S, Morganti, R, Jamrozy, M, Shulevski, A & Duncan, KJ 2020, 'Giant radio galaxies in the LOFAR Two-metre Sky Survey-I: Radio and environmental properties', *Astronomy and Astrophysics*, vol. 635, A5.
<https://doi.org/10.1051/0004-6361/201935589>

Digital Object Identifier (DOI):

[10.1051/0004-6361/201935589](https://doi.org/10.1051/0004-6361/201935589)

Link:

[Link to publication record in Edinburgh Research Explorer](#)

Document Version:

Peer reviewed version

Published In:

Astronomy and Astrophysics

General rights

Copyright for the publications made accessible via the Edinburgh Research Explorer is retained by the author(s) and / or other copyright owners and it is a condition of accessing these publications that users recognise and abide by the legal requirements associated with these rights.

Take down policy

The University of Edinburgh has made every reasonable effort to ensure that Edinburgh Research Explorer content complies with UK legislation. If you believe that the public display of this file breaches copyright please contact openaccess@ed.ac.uk providing details, and we will remove access to the work immediately and investigate your claim.



Giant radio galaxies in the LOFAR Two-metre Sky Survey-I

P. Dabhade,^{1,2*}, H. J. A. Röttgering¹, J. Bagchi², T. W. Shimwell^{1,3}, M. J. Hardcastle⁴, S. Sankhyayan⁵, R. Morganti^{3,6}, M. Jamroz⁷, A. Shulevski⁸, and K. J. Duncan¹

¹Leiden Observatory, Leiden University, P.O. Box 9513, NL-2300 RA, Leiden, The Netherlands

²Inter University Centre for Astronomy and Astrophysics (IUCAA), Pune 411007, India.

³ASTRON, The Netherlands Institute for Radio Astronomy, Postbus 2, 7990 AA, Dwingeloo, The Netherlands

⁴Centre for Astrophysics Research, School of Physics, Astronomy and Mathematics, University of Hertfordshire, College Lane, Hatfield AL10 9AB, UK

⁵Indian Institute of Science Education and Research (IISER), Dr. Homi Bhabha Road, Pashan, Pune 411008, India

⁶Kapteyn Astronomical Institute, University of Groningen, PO Box 800, 9700 AV, Groningen, The Netherlands

⁷Astronomical Observatory, Jagiellonian University, ul. Orla 171, 30-244 Kraków, Poland

⁸Anton Pannekoek Institute for Astronomy, University of Amsterdam, Postbus 94249, 1090 GE Amsterdam, The Netherlands

December 5, 2019

ABSTRACT

Giant radio galaxies (GRGs) are a subclass of radio galaxies which have grown to megaparsec scales. GRGs are much rarer than normal sized radio galaxies (< 0.7 Mpc) and the reason for their gigantic sizes is still debated. Here, we report the biggest sample of GRGs identified to date. These objects were found in the LOFAR Two-metre Sky Survey (LoTSS) first data release images, which cover a 424 deg^2 region. Of the 239 GRGs found, 225 are new discoveries. The GRGs in our sample have sizes ranging from 0.7 to 3.5 Mpc and have redshifts (z) between 0.1 and 2.3. Seven GRGs have sizes above 2 Mpc and one has a size of ~ 3.5 Mpc. The sample contains 40 GRGs hosted by spectroscopically confirmed quasars. Here, we present the search techniques employed and the resulting catalogue of the newly discovered large sample of GRGs along with their radio properties. We, here also show for the first time that the spectral index of GRGs is similar to that of normal sized radio galaxies, indicating that most of the GRG population is not dead or is not like remnant type radio galaxy. We find 20/239 GRGs in our sample are located at the centres of clusters and present our analysis on their cluster environment and radio morphology.

Key words. galaxies: jets – galaxies: active – radio continuum: galaxies – quasars: general

1. Introduction

A radio galaxy normally contains a radio core, jets and lobes powered by an active galactic nucleus (AGN). A radio galaxy that has grown to Mpc scales is traditionally defined as a giant radio galaxy (GRG) (Willis et al. 1974; Ishwara-Chandra & Saikia 1999). Here, the total size is defined as the largest angular separation between the end of the two radio lobes. This subclass of radio galaxies is among the largest single structures known in the universe along with the cluster radio relics (Röttgering et al. 1997; Bagchi et al. 2006; van Weeren et al. 2011). Born in the active nucleus of a galaxy or a quasar, radio galaxies/quasars eject collimated and bipolar relativistic jets (Lynden-Bell 1969; Begelman et al. 1984). The driving engine for these jets is an accreting super massive black hole (SMBH) with a typical mass of $10^8 - 10^{10} M_\odot$. SMBHs that drive powerful jets reside in elliptical galaxies and only a handful are found in spiral galaxies (Hota et al. 2011; Bagchi et al. 2014).

Morphologically, radio galaxies have been historically divided into two classes, Fanaroff-Riley type I (FR-I) and Fanaroff-Riley type II (FR-II). The lower radio luminosity FR-I sources have their brightest regions closer to the nucleus and their jets fade with distance from core. For the

more powerful Fanaroff-Riley type II (FR-II) radio galaxies (Fanaroff & Riley 1974), the jet remains relativistic all the way from the central AGN to the hotspots in lobes.

For all radio galaxies the ejection of a collimated jet is dependent on the availability of fuel. Assuming GRGs grow to an enormous size due to a prolonged period activity, this would require either an unusually large reservoir of fuel or a very efficient jet formation mechanism.

Radio galaxies were first discovered about six decades ago (Jennison & Das Gupta 1953) and since then hundreds of thousands of radio galaxies have been found. In contrast to the large number of radio galaxies, only ~ 350 (Dabhade et al. (2017); Kuźmicz et al. (2018) & references therein) or so GRGs have been found of which only a small fraction has been studied in detail. These giants, when associated with quasars as their AGN, are called giant radio quasars (GRQs) and only around 70 GRQs are known so far (Kuźmicz et al. 2018). The term ‘GRQ’ is used here to emphasize the fact that the GRG has quasar as the powering AGN.

The hypotheses proposed (which are not mutually exclusive) to explain the enormous sizes of GRGs include:

1. GRGs possess exceptionally powerful radio jets when compared to normal radio galaxies and these provide

* E-mail: pratik@strw.leidenuniv.nl

the necessary thrust to reach Mpc scales (Wiita et al. 1989).

2. GRGs are very old radio galaxies and have had sufficient time to expand over large distances (Subrahmanyam et al. 1996).
3. GRGs grow in low density environments, (Mack et al. 1998; Malarecki et al. 2015; Saripalli & Malarecki 2015) enabling them to grow comparatively fast.

None of the above hypotheses have been tested using large uniform samples of GRGs. In smaller samples, contradictory results have been found. For example, Mack et al. (1998) found evidence that the ages of GRGs in their sample are similar to that of normal sized radio galaxies which is contradictory to the above mentioned second point. Also, point three has been contradicted by findings of Komberg, & Pashchenko 2009; Dabhade et al. 2017, where they have reported a number of GRGs to be located in cluster environments.

The exceptionally large lobes of GRGs makes them excellent laboratories for studying the evolution of the particle and magnetic field energy density, acceleration of high energy cosmic rays and can also be used to probe large scale environments (Kronberg et al. 2004; Safouris et al. 2009; Malarecki et al. 2015; Isobe, & Koyama 2015).

Until the mid 1990s, GRGs were mostly discovered serendipitously. Only after the advent of deep and large sky radio surveys like Faint Images of the Radio Sky at Twenty-cm (FIRST) (Becker et al. 1995), NRAO VLA Sky Survey (NVSS) (Condon et al. 1998), Westerbork Northern Sky Survey (WENSS) (Rengelink et al. 1997) and Sydney University Molonglo Sky Survey (SUMSS) (Bock et al. 1999) were systematic searches for GRGs carried out. Lara et al. (2001), Machalski et al. (2001) and Dabhade et al. (2017) searched the higher frequency survey such as NVSS (1400 MHz) for GRGs, and Saripalli et al. (2005) used SUMSS, whereas Schoenmakers et al. (2001) used lower frequency surveys such as WENSS (327 MHz).

The lobes of the GRGs have steep spectral indices and hence are bright at low radio frequencies. Cotter et al. (1996) made a sample of GRGs using the 151 MHz 7C survey (McGilchrist & Riley 1990). The 7C survey has a low resolution of 70×70 cosec (δ) arcsec² and a noise level ~ 15 mJy beam⁻¹ (1σ) and as a consequence contained only a few GRGs. The WENSS has better resolution (54×54 cosec (δ) arcsec²) and better sensitivity (RMS noise (1σ) ~ 3 mJy beam⁻¹) and also covers a somewhat larger area ~ 8100 deg² (7C survey ~ 5580 deg²). This enabled Schoenmakers et al. (2001) to compile a large sample of 47 GRGs from the WENSS.

In recent years, four large low frequency surveys have been carried out, namely :

- 119-158 MHz Multifrequency Snapshot Sky Survey (MSSS) (Heald et al. 2015).
- 150 MHz TIFR GMRT SKY SURVEY- Alternative data release-1 (TGSS-ADR1) (Intema et al. 2017).
- 72-231 MHz GaLactic and Extragalactic All-sky Murchison Widefield Array (GLEAM) survey (Hurley-Walker et al. 2017).
- 120 - 168 MHz LOFAR Two-metre Sky Survey (LoTSS) (Shimwell et al. 2017, 2019).

In the past 20 years, large surveys have also been carried out at optical wavelengths. These surveys include the

Sloan Digital Sky Survey (SDSS) (York et al. 2000), the 2 degree Field Galaxy Redshift Survey (2dFGRS) (Colless et al. 2001), the 2MASS Redshift Survey (2MRS) (Huchra et al. 2012), the 6 degree Field Galaxy Survey (6dFGS) (Jones et al. 2009) and most recently the deep photometric survey called Panoramic Survey Telescope and Rapid Response System (Pan-STARRS; Kaiser et al. 2002, 2010; Chambers et al. 2016). The data from these surveys has allowed the identification of many new GRGs as shown in Dabhade et al. (2017).

High sensitivity to low surface brightness features and high spatial resolution to decipher the morphologies are key requirements in identifying GRGs. LoTSS provides combination of both these properties for the first time and hence, combining it with the SDSS/Pan-STARRS optical surveys, we use it to search for new GRGs in order to form a statistically significant sample. Our study of GRGs will be presented in 2 papers:

1. Paper I (this paper) reports the methodology used for the systematic search scheme implemented for the discovery of new GRGs/GRQs from the LoTSS and presents the sample's radio properties.
2. Paper II will focus on studying the host AGN and galaxy properties of the GRGs/GRQs sample and comparing them with another sample (also from LoTSS) of normal sized radio galaxies (NRGs) matched in redshift and optical/radio luminosity to the GRG sample.

Throughout the paper, we adopt the flat Λ CDM cosmological model based on the latest Planck results ($H_o = 67.8$ km s⁻¹ Mpc⁻¹, $\Omega_m = 0.308$) (Planck Collaboration et al. 2016), which gives a scale of 4.6 kpc/" for the redshift of 0.3. In this paper, radio galaxies with projected linear sizes ≥ 0.7 Mpc, computed using the above mentioned cosmological parameters are called GRGs. All images are in the J2000 coordinate system. We use the convention $S_\nu \propto \nu^{-\alpha}$, where S_ν is flux at frequency ν and α is the spectral index.

2. Identifying new GRGs in LoTSS

2.1. The LoTSS first data release

LoTSS is a 120-168 MHz survey that is being conducted with the high-band antennas (HBA) of LOFAR and will eventually cover the whole northern sky. Hardcastle et al. (2016) have already demonstrated the potential of LoTSS deep observations for discovering GRGs and found seven in the *Herschel* ATLAS North Galactic Pole survey area (142 deg²). Here, we focus on the LoTSS first data release (LoTSS DR1; Shimwell et al. 2019). The LoTSS DR1 spans (J2000.0 epoch) right ascension 10h45m to 15h30m and declination 45°00' to 57°00' (HETDEX:Hobby-Eberly Telescope Dark Energy Experiment Spring field region) covering an area of 424 deg² with a median noise level across the mosaic of 71μ Jy beam⁻¹ and $\sim 6''$ resolution. In Fig. 1, we see a comparison between the LoTSS and other radio surveys like NVSS and FIRST for a GRG from our sample. The top image (a) is an optical-radio overlay with blue colour indicating LoTSS low frequency $6''$ resolution map on the optical SDSS tri-colour image. In the bottom image (b), the same source is shown as it is observed in NVSS and FIRST. NVSS, though highly sensitive to large scale diffuse emission, fails to reveal the finer details across the source

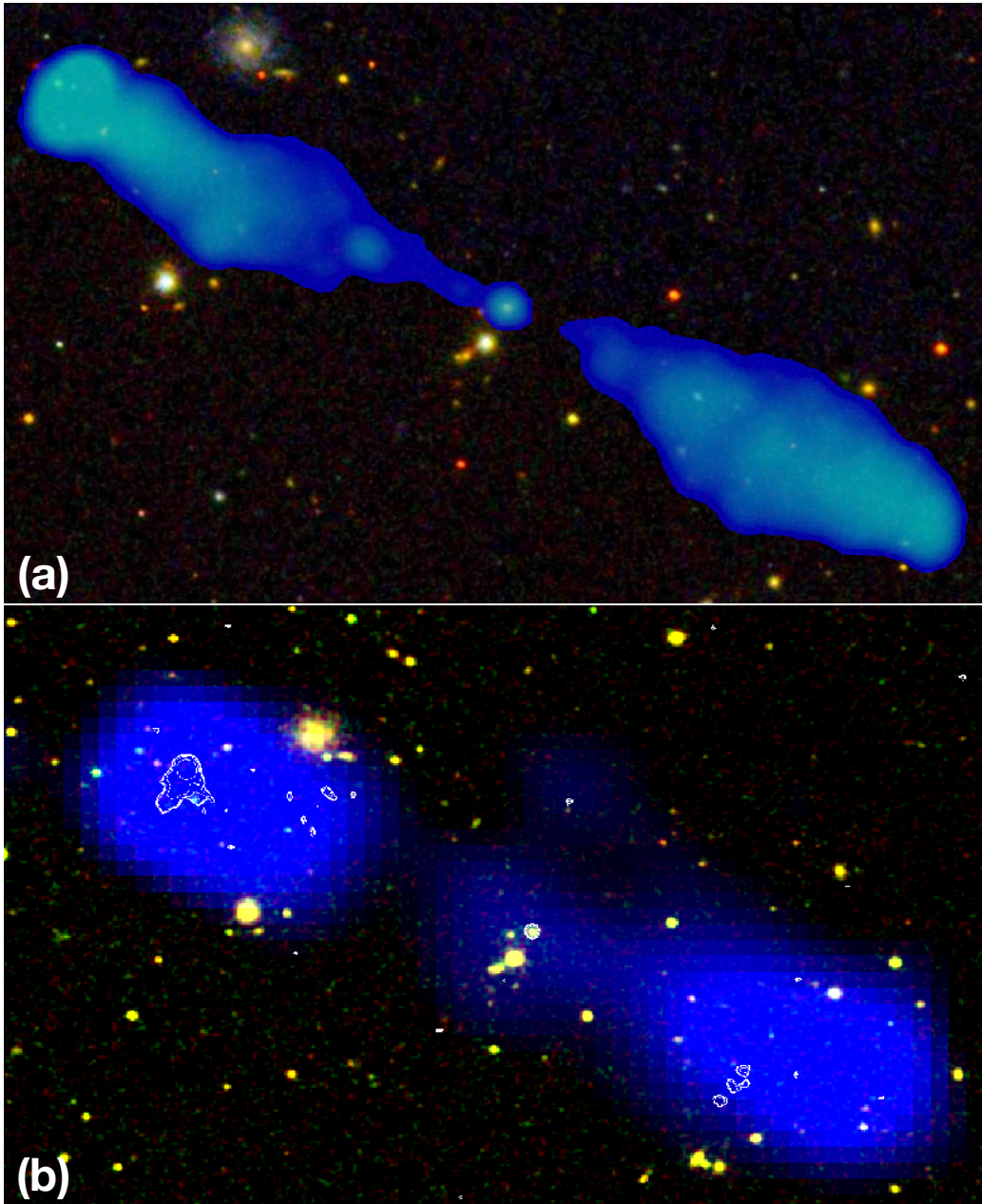


Fig. 1: (a) A colour composite image of 1.86 Mpc long GRG J105817.90+514017.70 made using LoTSS-DR1 144 MHz radio and optical SDSS image. (b) A colour composite image of the same object with optical and radio overlay, where blue represents 1400 MHz NVSS image with 45'' resolution and white contours from 1400 MHz FIRST survey having 5'' resolution superimposed on optical SDSS image. The LoTSS 144 MHz 6'' resolution image shown in top panel clearly resolves the core and jet and also highlights the diffuse parts of the lobes, which is missed by the FIRST and unresolved in NVSS.

and cannot properly resolve the core due to its coarser resolution of 45''. FIRST survey on the other hand has high resolution, which manages to resolve the core and hence help in identifying the host AGN/galaxy but misses out on almost all the diffuse emission of the lobes and hence it alone cannot be used to identify RGs. The LoTSS data provides with both high resolution as well as high sensitivity and does not resolve out structures revealing finer details of the emission of large scale jets. Both images (Fig. 1 and source 7 of Fig. A.1) show clearly the radio core and jets feeding the giant radio lobes and the hotspots. This clearly

illustrates the excellence of LoTSS and its great potential in unveiling interesting sources.

Using the LoTSS DR1 radio data and optical-infrared data, a Value Added Catalogue¹ (VAC) of 318520 radio sources has been created (Williams et al. (2019); DR1-II). The host galaxies/quasars were identified using the Pan-STARRS and Wide-field Infrared Survey Explorer (WISE; (Wright et al. 2010)). The Pan-STARRS-AllWISE catalogue was cross matched with LoTSS survey using a likelihood ratio method. Furthermore, human visual inspection

¹ <https://www.lofar-surveys.org/>

was used for the final classification of complex radio sources using the LOFAR Galaxy Zoo (LGZ), the details of which are given in Williams et al. (2019) (DR1-II). The photometric redshift and rest-frame colour estimates for all hosts (galaxies/quasars) of the matched radio sources are presented in Duncan et al. (2019) (DR1-III). The VAC lists the radio properties, identification methods and optical properties where available.

2.2. Semi-automated search for GRGs

The methodology of identifying GRGs and forming the final catalogue is presented in flow chart as seen in Fig. 2 and is described below:

1. The VAC of 318520 radio sources was at first refined based on the point source completeness, which is 90% at an integrated flux density of 0.5 mJy for LoTSS DR1 (Shimwell et al. 2019). We apply a flux density cut at this level. This results in the total number of sources reducing to 239845.
2. Secondly, only objects with optical identification and redshift estimates were selected resulting in 162249 sources.
3. The redshift information in the VAC is compiled mainly using the SDSS spectroscopic data. For the sources which do not have spectroscopic redshift information from the SDSS or any other spectroscopic survey, Duncan et al. (2019) have estimated photometric redshifts using multi-band photometry. For our work, we have imposed a further photometric quality cut on the estimated photometric redshifts. Only sources satisfying the condition $\Delta z/(1+z) < 0.1$ were selected, where Δz is the half width of the 80% credible interval. This results in reduction of the sample from 162249 sources to 89671 sources.
4. Williams et al. (2019) inspected all the extended complex radio sources visually via the LGZ program and estimated the angular size for them. A semi-automated way was adopted for this work by Williams et al. (2019). This resulted in angular size estimation of total 13222 sources in the VAC.

We find that, of the 89671 sources (which have reliable redshifts), only 4808 sources have angular size estimates by the above mentioned method (LGZ). Or, in other words, 4808 sources have angular size estimates as well have passed the photometric quality cut (point 3).

5. For all the 4808 sources, we computed the projected linear size (kpc) and only those that had an extension above 700 kpc were considered for further analysis. This resulted in a sample of 398 candidate GRGs.
6. Further, the candidate GRGs were visually inspected by us to identify and remove those with uncertainty in the host, large asymmetry or a high degree of bending or narrow angle tailed morphology. The angular size measurement was refined by taking the distance between the farthest points of the 3σ contours of the source and the projected linear size was recomputed. We have measured the sizes of the sources using the LoTSS DR1- low resolution maps (20'') as it shows the entire structure of the GRGs and also does not miss any diffuse emission. The final sample size of GRGs from the VAC is 186 GRGs.

2.3. Manual Visual Search from LoTSS DR1

An independent manual visual search (MVS) was also carried out to search for GRGs in LoTSS DR1 to look for additional GRGs that were missed in the semi-automated method employed on the VAC. In this approach, all mosaics from LoTSS DR1 were scanned for extended double lobed structures (candidate GRGs). The radio core of the candidate GRGs were searched in the optical band (SDSS and PAN-STARRS) and IR band (WISE) for counterparts (host galaxies). We also used other radio surveys like FIRST, TGSS and WENSS for additional information and consistency checks. This method (MVS) yielded 53 additional GRGs. These 53 GRGs were missed by the selection criteria because of the photometric redshift cuts we imposed or because the complex nature of the source structures lead them to being not fully characterised in the VAC.

2.4. Final catalogue: VAC+MVS

We combined both GRG samples (VAC and MVS) to form the final GRG catalogue of 239 GRGs (Table. 2). The final catalogue of 239 GRGs was cross matched with the GRG catalog of Kuźmicz et al. (2018), which is a complete compendium of GRGs published till 2018 along with other literature search, and we find 13 of our 239 GRGs to be already known (listed in 13th column of Table 2). The high-resolution 6'' LoTSS images at 144 MHz of GRGs can be found in the appendix A from Fig. A.1 to Fig. A.8, where the white contours represent emission seen in the LoTSS low resolution 20'' map at 144 MHz. The position of host galaxy is marked with a white cross '+'.

3. Results : The LoTSS catalogue of GRGs

Our search of LoTSS DR1 has enabled us to construct a catalog² of 239 GRGs. With the high sensitivity of LoTSS, we are able to detect GRGs as faint as ~ 2.5 mJy in total flux at 144 MHz. Using the available optical data and radio data, we have computed the radio powers, spectral indices and classified the morphological types for the sample of 239 GRGs (Table. 2). The GRGs from our sample have sizes in the range of 0.7 Mpc to ~ 3.5 Mpc (Fig. 3) with median size of 0.89 Mpc and mean size of 1.02 Mpc. GRGs with sizes greater than 2 Mpc are very rare within the GRG population and in our sample we find 7/239 GRGs having sizes ≥ 2 Mpc.

3.1. Optical host properties

The GRGs in our sample span a wide redshift range from 0.1 to 2.3 (Fig. 4). GRGs that are hosted by galaxies are not selected beyond $z \sim 1$ due to sensitivity limits of the optical surveys. We use the SDSS DR14 Quasar catalog (Pâris et al. 2018) and SDSS to identify the GRGs hosted by quasars in our sample. 151/239 GRGs have optical spectroscopic

² Before the precise measurement of H_0 , originally GRGs were defined as RGs with projected linear sizes ≥ 1 Mpc using $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$. If we convert the original definition with the latest precise measurement of $H_0 = 67.8 \text{ km s}^{-1} \text{ Mpc}^{-1}$ with $\Omega_m = 0.308$, then the lower limit of projected linear size of GRGs is ~ 0.74 Mpc. In our sample, 28 GRGs have projected linear sizes between 0.7 Mpc to 0.74 Mpc.

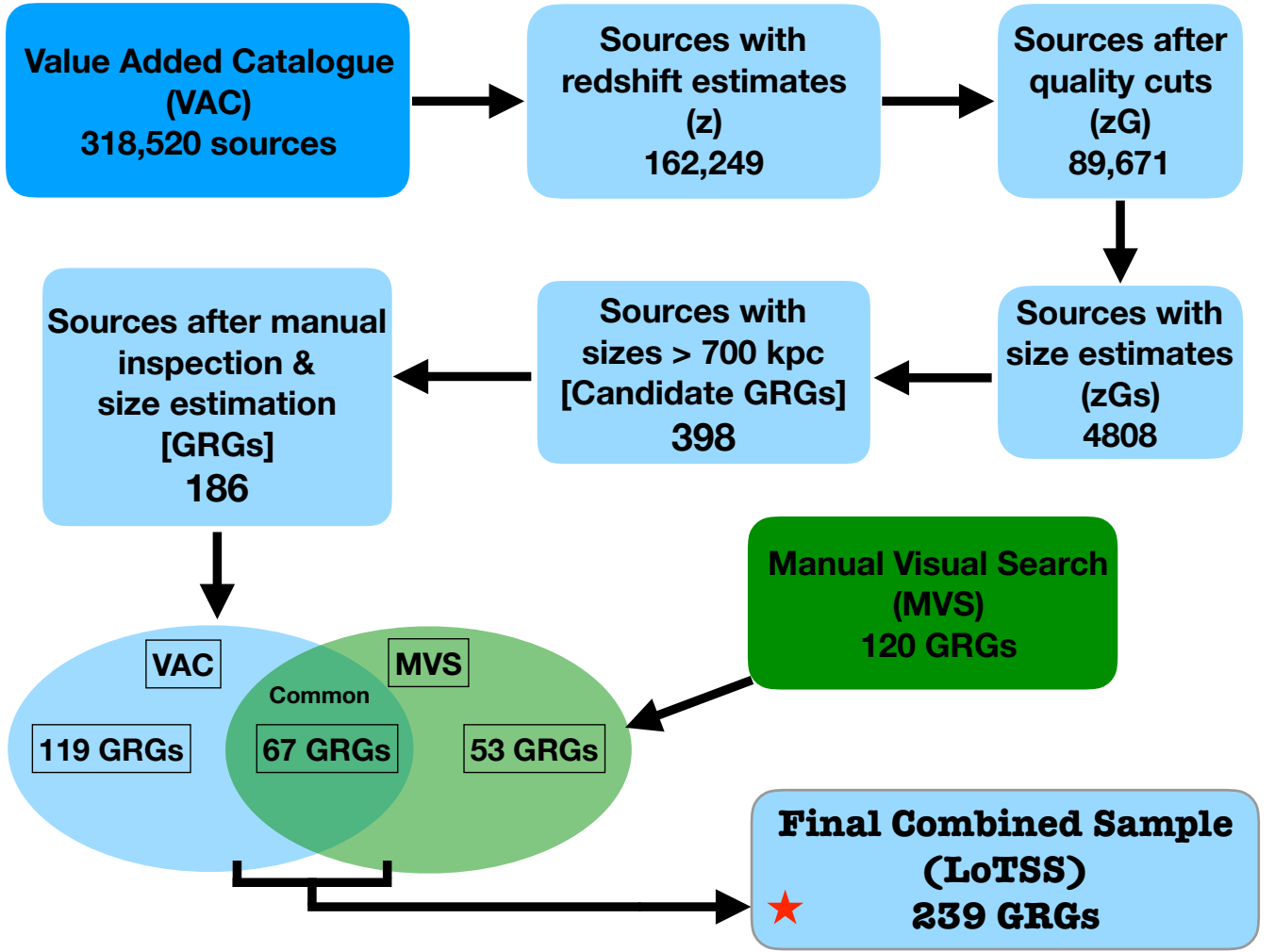


Fig. 2: The above figure shows schematics for finding GRGs from LoTSS DR1 using VAC and MVS in steps. More details are presented in Sect. 2.

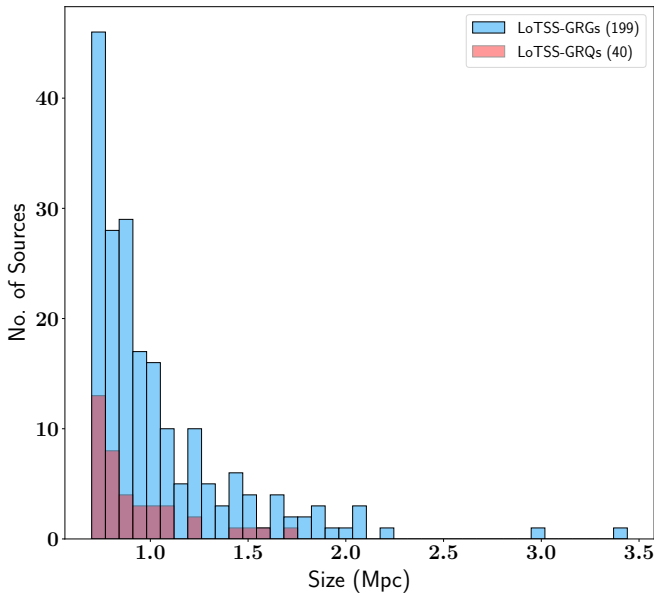


Fig. 3: Histogram for sizes of GRGs and GRQs from the LoTSS sample.

redshifts from SDSS. Of the 151 GRGs with optical spectroscopic redshifts 40 are hosted by quasars (GRQs). Interestingly, based on the available optical data (SDSS and Pan-STARRS) and the Galaxy Zoo catalog of spiral galaxies (Hart et al. 2016), we find that none of the GRGs is hosted by a spiral galaxy.

Most of the redshifts of the hosts of GRGs were obtained from the VAC, which has made use of the spectroscopic data from the SDSS and has estimated photometric redshifts for the sources which do not have spectroscopic data in SDSS or in other literature. The uncertainties in the photometric redshifts are explained in section 3 (especially section 3.5) of Duncan et al. 2019.

3.2. Spectral Index (α_{144}^{1400})

The LoTSS DR1 provides maps of the radio sky (centered at 144 MHz) at two resolutions, high (6'') and low (20'') which enables us to identify compact cores, jets and hotspots as well as see the extent of diffuse emission. To compute the flux at 144 MHz, we use the 20'' low resolution maps of LoTSS DR 1 (column 8 of Table 2). Measurement in the flux errors is done taking a 20% calibration error for LoTSS

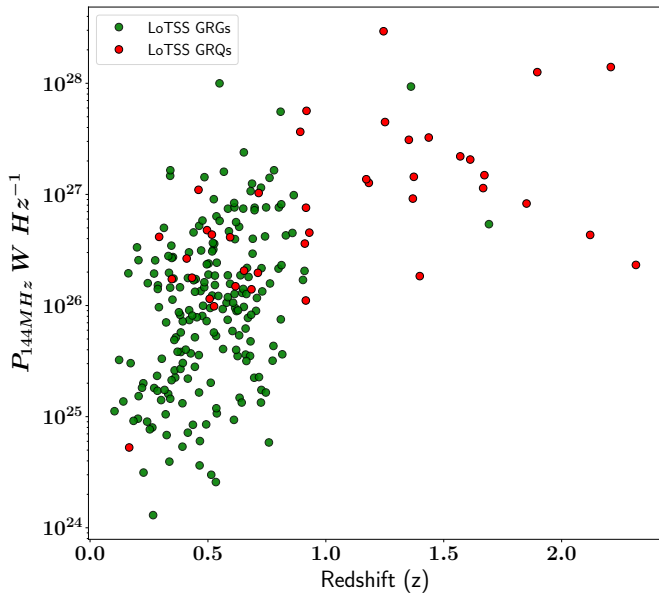


Fig. 4: Radio power and redshift distribution for GRGs and GRQs from LoTSS. The two galaxies marked in green colour with $z \geq 1$ are quasar candidates as explained in Sec. 3.4.

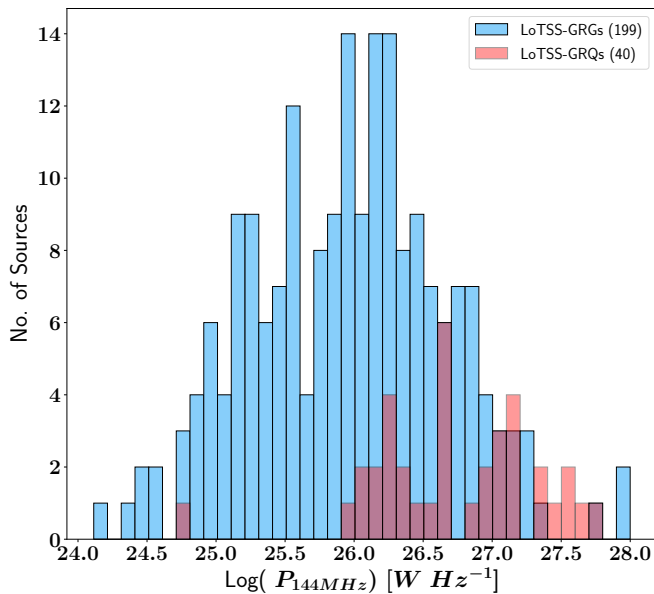


Fig. 5: Histogram of radio power at 144 MHz of GRGs and GRQs.

DR1 (Shimwell et al. 2019). By combining high frequency (1400 MHz) NVSS and low frequency (144 MHz) LoTSS, we have computed the integrated spectral index (α_{144}^{1400}) for GRGs in our sample (column 10 of Table 2). The NVSS radio map cutouts for GRGs were obtained from its server³, from which we have measured the flux of the source.

Following steps were adopted to obtain spectral index:

- Convolve the LoTSS low resolution images (cutouts of GRGs from main mosaics) to the same resolution of

NVSS (45'') and regrid the LoTSS images to match with NVSS.

- Make automated masks (regions to extract flux) using the package PyBDSF⁴ (Python Blob Detection and Source Finder) of Mohan & Rafferty (2015).
- The 45'' convolved LoTSS maps were manually inspected for possible contamination (using the high resolution LoTSS maps (6'') and FIRST's 5'' maps) from other sources in the field and manual masks were made. The flux was obtained by considering only the region in manual masks from the automated masks.
- Finally, using the fluxes obtained from NVSS and LoTSS, integrated spectral index was computed for the whole source.

A total of 37/239 sources were contaminated by other nearby radio sources in the low resolution convolved LoTSS maps and NVSS maps, for these sources (marked with '-' in column 10 of Table 2) a spectral index was not computed. For sources with no detection in NVSS, upper limits on the flux were computed and a spectral index limit was obtained (indicated with < sign in column 10 of Table 2).

Fig. 6 shows the spectral index (α_{144}^{1400}) distribution of 171 GRGs and 31 GRQs. The median and mean values for spectral index of GRGs are 0.77 and 0.79, respectively. Similarly for GRQs, the median and mean values are 0.78 and 0.76, respectively. These mean spectral index values of GRGs and GRQs are similar to those of normal sized radio galaxies (Oort et al. 1988; Gruppioni et al. 1997; Kapahi et al. 1998; Ishwara-Chandra et al. 2010; Mahony et al. 2016). In the work of Oort et al. (1988), they have surveyed the Lynx field with the Westerbork Synthesis Radio Telescope (WSRT) at 325 MHz and 1400 MHz. Gruppioni et al. (1997) surveyed the Marano field using the Australia Telescope Compact Array (ATCA) at 1.4 GHz and 2.4 GHz. Kapahi et al. (1998) surveyed the Molonglo Radio Catalogue sources with the VLA at L and S bands. Ishwara-Chandra et al. (2010) surveyed the LBDS-Lynx field (LBDS: Leiden-Berkeley Deep Survey) using the 150-MHz band of Giant Metrewave Radio Telescope and other archival data from GMRT at 610 MHz and 325 MHz along with data from other surveys like WENSS, FIRST and NVSS. Mahony et al. (2016) surveyed the Lockman Hole field using LOFAR 150 MHz and WSRT 1.4 GHz. All the above work obtained spectral index values in the range of ~ 0.7 to 0.8 and therefore 0.75 value is often assumed as the average spectral index value for radio galaxies in absence of any multi-frequency observations.

The above mentioned result implies that RGs and GRGs do not differ in terms of their spectral index properties and most of the GRG population are active radio galaxies and not dead or remnant radio galaxies. The values for spectral index of GRGs given by us are integrated values for the whole source and not just one component of GRGs (like core or lobes). It is found that for some GRGs, the lobes come out to be more steeper than usual- which is obtained by multi-frequency (observations at 3 to 4 radio bands) observations of individual sources. There are examples (source number 16,17 and 193) in our sample where the sources are core dominated (maximum flux observed in radio core) and it hence influences sources to have flatter (less than 0.5) overall spectral index.

³ <https://www.cv.nrao.edu/nvss/postage.shtml>

⁴ <http://www.astron.nl/citt/pybdsf/>

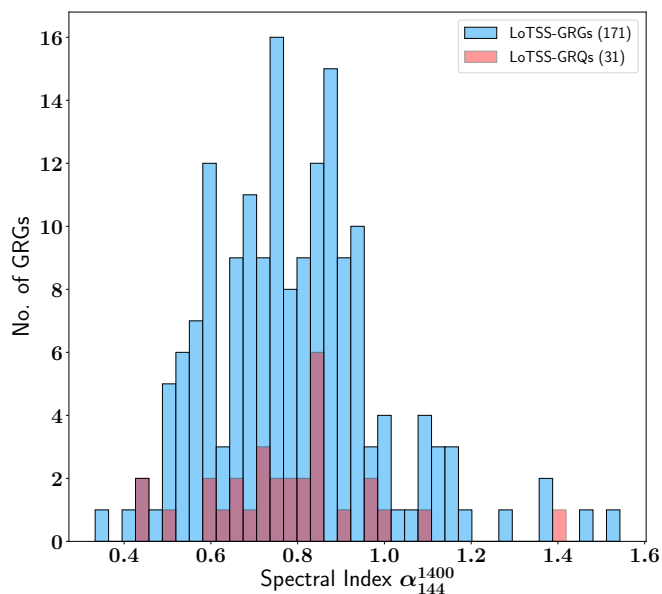


Fig. 6: Histogram of Spectral index (α_{144}^{1400}) of GRGs and GRQs in LoTSS DR1 sample made using LoTSS and NVSS.

3.3. Notes on individual objects

Below, we present brief notes on a selection of some of the most interesting objects in our sample.

- Source 5- GRG J105725.96+492900.31: The host galaxy does not perfectly coincide with the bright core like feature which are separated by $\sim 5''$ distance from one another. It is most likely a knot in the jet, which is located right next to a faint radio core. There is no other galaxy present in the immediate vicinity and no other alternate compact radio emission along with axis of the source, which could possibly be an alternate host of this source.
- Source 10- GRG J110433.11+464225.76: This unusual source has a well detected radio core but does not exhibit any jets or well formed lobes. It has diffuse emission on either side of the radio core and has been reported by [Thwala et al. \(2019\)](#) as a relic radio galaxy with a size of 0.86 Mpc, thereby making it a GRG. Towards north western side of the source, there exists a independent FR-II type radio galaxy. Using the higher resolution maps the independent source's flux was measured and then subtracted from the LoTSS's low resolution maps to avoid contamination of source fluxes.
- Source 44- GRG J113931.77+472124.3: As seen in Fig. A.2, there is compact radio source north of the '+' marker (host galaxy-radio core), which is an independent source. Also, the bright compact source seen towards south eastern side is independent of the GRG J113931.77+472124.3.
- Source 61- GRG J121555.53+512416.41 and Source 162- GRG J135628.50+524219.23: These sources are core dominated objects at both low and high frequencies. The lobes of both sources are not detected in NVSS. These are possibly candidates of revived/rejuvenated GRGs as we only observe the diffuse lobes at low frequencies (LoTSS), which could be from the previous epoch. The presence of the bright core may indicate restarting activity.

- Source 136- GRG J133322.79+533250.94: This source displays a peculiar morphology and is possibly residing in an unique environment. The VAC estimates its photometric redshift (z) to be 0.3539 ± 0.0344 . [Lopes \(2007\)](#) estimates its photometric redshift based on SDSS data to be 0.39301 ± 0.02578 . GRG J133322.79+533250.94 is close to two galaxy clusters, namely WHL J133322.0+5333490 at redshift of 0.3938 and WHL J133316.5+533333 at a redshift of 0.3834. Based on [Lopes \(2007\)](#) photometric redshift, GRG J133322.79+533250.94 is plausibly to be associated with the galaxy cluster WHL J133322.0+5333490 as they are at similar redshifts and is separated by $\sim 1'$. This GRG appears to be residing in an environment of a possible merger of two galaxy clusters and the associated radio relic can be seen towards north of the GRG (Fig. 7).
- Source 161- GRG J135414.72+491315.2: The radio core coinciding with a faint galaxy is detected only in the FIRST survey.
- Source 221- GRG J145002.36+540528.27: The radio core is well detected in the FIRST survey, which coincides with a galaxy. The source is sufficiently resolved only in the LoTSS high resolution map as it is located right next to a bright source. The spectral index was computed by estimating an upper limit for the flux from NVSS.
- Source 222- GRG J145057.28+530007.76: The radio core is only resolved in the FIRST survey. It is not symmetrically placed between the two lobes and is closer to the eastern lobe.
- Source 237- GRG J151835.37+510410.70: The radio core is not symmetrically placed between the two lobes and is closer to the northern lobe. It is also well detected in the FIRST survey. The southern lobe shows a prominent hotspot.

3.4. Giant radio quasars

GRGs with quasars as their AGN (GRQs) are even more rarer than the GRGs with non quasar AGN. Based on the catalogue of [Kuźmicz et al. \(2018\)](#) and candidates presented by [Kuligowska, & Kuźmicz \(2018\)](#), there are only about ~ 70 GRQs known till date and less than ~ 10 GRQs known at $z \geq 2$. From our LoTSS sample of 239 GRGs, 40 are confirmed GRQs and 2 are candidates GRQs (Q_c). A total of only six GRQs of the 40 confirmed GRQs were previously reported (column 13 in Table 2) and rest are all new findings. Our sample of ~ 40 GRQs significantly increase the sample of total GRQs known till date from 70 ([Kuźmicz et al. 2018](#)) to more than 100. The GRQs in our sample have median and mean radio powers of $6.2 \times 10^{26} \text{ W Hz}^{-1}$ and $2.4 \times 10^{27} \text{ W Hz}^{-1}$ at 144 MHz respectively as seen in the histogram distribution of radio powers (Fig. 5).

Below we present some brief notes on individual objects-

- GRQs J110833.99+483203.03, J123933.21+500708.01 and J133418.63+481317.08 are the three GRQs which we have found at very high redshift of $z \geq 2$ and are shown in Fig. 8.
- GRG J110833.99+483203.03 (16 in Table 2) is a highly core dominated object at both high as well as low frequencies and hence exhibits a very flat spectral index of ~ 0.2 .

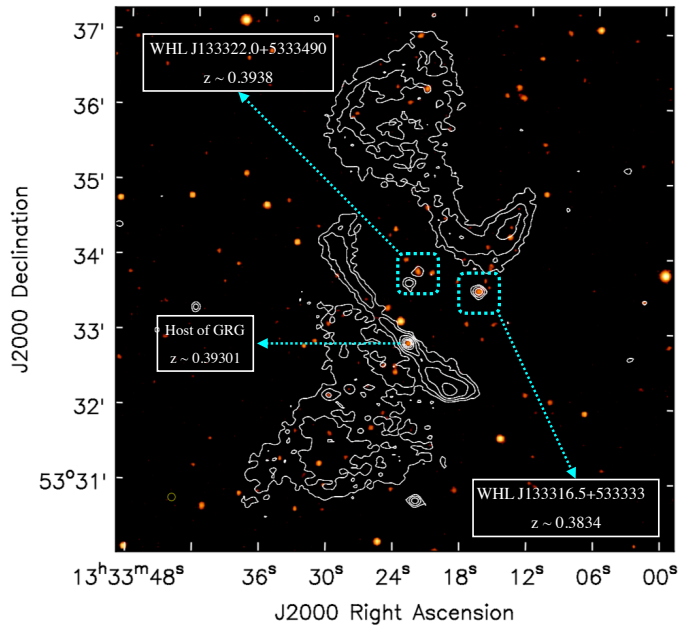


Fig. 7: The figure shows GRG J13322.79+533250.94 amidst possibly galaxy cluster relics. The background image in orange colour is SDSS I band optical image which is superimposed with LoTSS DR1 high resolution ($6''$) contours.

- GRGs J132554.31+551936.23 and J141408.45+484156.11 (126 and 178 in Table 2) which have high redshifts and are the two quasar candidates (Q_c) as indicated by Richards et al. 2009, 2015. These objects have all the properties of a radio loud quasars but still are candidates because of absence of an optical spectra. Therefore they are not labeled as quasars (not marked in red colour) in any of the figures.

3.5. Environment analysis of GRGs

We use two optically selected galaxy cluster catalogs to identify host of GRGs with Brightest Cluster Galaxies (BCGs), which are found to be at the centers of galaxy clusters. We chose the following two catalogues as they are made using SDSS and have an overlap with HETDEX region which is common for LoTSS DR1.

Firstly, we used the galaxy cluster catalog of Wen et al. (2012) (here after WHL cluster catalog), which consists of 132,684 clusters. They have used photometric data from SDSS-III to find the clusters. This is the biggest galaxy cluster catalog made using SDSS in the redshift range of $0.05 < z < 0.8$ and is $\sim 95\%$ complete for clusters with a mass of $M_{200} > 10^{14} M_\odot$ in the redshift range of $0.05 < z < 0.42$. A total of 71 GRGs from our sample fall in this redshift range.

We find 17 GRGs to be BCGs based on the WHL cluster catalog. We also used the Gaussian Mixture Brightest Cluster Galaxy (GMBCG) catalogue (Hao et al. 2010) consisting of 55,880 galaxy clusters which is more than 90% complete within the redshift range of $0.1 < z < 0.55$. We find 3 extra GRGs to be BCGs from the GMBCG catalogue. Therefore, we find a total of 20 GRGs (Table 3) from our sample of 239 to be BCGs residing in dense cluster environ-

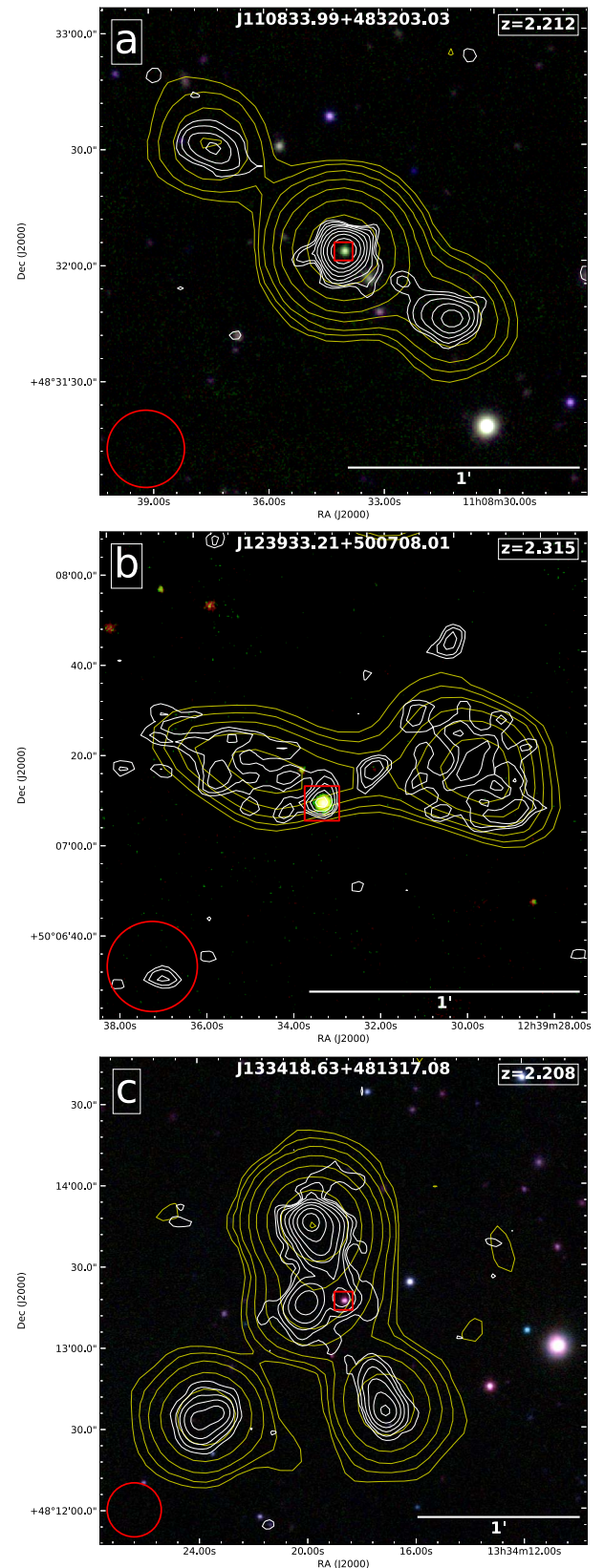


Fig. 8: The above three images are optical-radio overlay of high redshift ($z > 2$) GRGs, where the optical images composite of three bands of PanSTARRS and radio is from the LoTSS DR-1 at two resolutions. The red square indicates the host of the three GRGs. The white and yellow colour contours represents LoTSS high ($6'' \times 6''$) and low ($20'' \times 20''$; red circles at bottom left) resolution maps respectively, which have eight equally spaced relative contour levels starting from three times σ , where σ denotes the local RMS noise.

ments and hosted by BCGs of the clusters. The mass and radius (obtained from Wen et al. 2012) of the 17 clusters are listed in Table 3.

4. Discussion

4.1. GRGs with sizes > 2 Mpc

GRGs with projected linear sizes > 2 Mpc are quite rare and as observed from Fig. 3, there are only 7 GRGs with sizes greater than 2 Mpc in our sample of 239. Based on sample compilation (incomplete) of all the known GRGs till date of ~ 400 sources, only ~ 62 GRGs ($\sim 15\%$) have sizes greater than 2 Mpc (Willis et al. 1974; Laing et al. 1983; Lacy et al. 1993; Ishwara-Chandra & Saikia 1999; Schoenmakers et al. 2001; Machalski et al. 2001; Lara et al. 2001; Saripalli et al. 2005; Machalski et al. 2007; Koziel-Wierzbowska & Stasińska 2011; Solovyov, & Verkhodanov 2014; Amirkhanyan et al. 2015; Amirkhanyan 2016; Dabhade et al. 2017; Clarke et al. 2017; Prescott et al. 2018; Sebastian et al. 2018).

If we take a complete radio sample like 3CRR (Laing et al. 1983), then we have only 12 GRGs from the total sample of 173, which is only $\sim 7\%$ of the total population. Considering this sample of 12 GRGs from 3CRR⁵, there are only 2 sources with sizes greater than 2 Mpc, which constitutes $\sim 17\%$ of the total GRG population in 3CRR complete sample. The conditions under which a radio galaxies grows to become a GRG is not very well understood and studies so far based on the observations and theory seems to suggest that the gigantic size could be due to an interplay between the jet power and the environment.

4.2. Morphology of GRGs

Most known GRGs exhibit FR-II type of morphology and we very rarely see radio galaxies with Mpc scale size with FR I type of morphology. Based on their morphologies we classify 18 GRGs as FR I type and 215 GRGs as FR II type. GRG J121900.76+505254.41 is the only FR-I GRG residing at cluster centre and hosted by a BCG (Table 3).

The total radio power of FR-I sources on average is less than that of the FR-II sources (Ledlow & Owen 1996; Lara et al. 2004). As seen in Fig. 9, the FR-I type of GRGs have a limited range in radio power $10^{24} \sim 10^{26} \text{ W Hz}^{-1}$, whereas the FR-II exhibits a wide range of radio powers from $10^{24} \sim 10^{28} \text{ W Hz}^{-1}$ at 144 MHz.

4.2.1. GRGs with Hybrid morphology

Radio galaxies which show FR-I morphology on one side and FR-II morphology on other side are referred as HyMoRS (Gopal-Krishna & Wiita 2002). The earliest example of such morphology was presented and studied by Saikia et al. (1996), for the radio galaxy 4C +63.07. They attribute this to an intrinsic asymmetry in either the collimation of its jets or the supply of fuel from the central black hole to opposite sides.

Gawroński et al. (2006) have estimated the occurrence of HyMoRS to be as low as $\leq 1\%$ amongst the radio galaxy population. Recently, Kapińska et al. (2017) presented 25 new candidate HyMoRS, of which 5 are GRGs with one

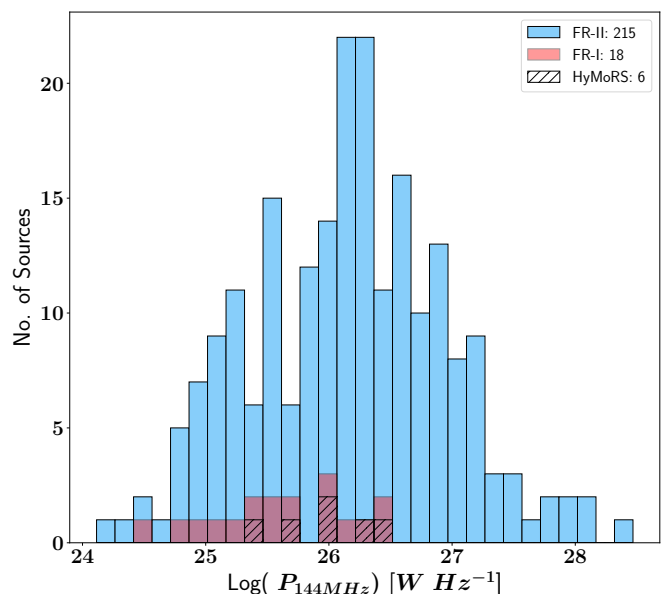


Fig. 9: Radio power (144 MHz) distribution of FR-I, FR-II and HyMoRS sources in LoTSS DR1 sample of giants.

being at the centre of a galaxy cluster. Therefore, one of the possible scenarios for such hybrid morphology could be attributed to different environments on each side of the host galaxy (radio core).

In our sample of 239 GRGs, we find 6 examples of HyMoRS (identified via visual inspection) which are listed in column 11 (FR type) of Table 2, indicated with roman numeral III, images of these sources are presented in Fig. 11. This is by far the largest HyMoRS GRG sample reported ever. Environment as well as host AGN studies are needed to understand this class of radio galaxies, which can also grow to megaparsec scales in size.

4.2.2. DDRGs: Double double radio galaxies

DDRG are FR-II type objects with two pairs of lobes, which is indicative of their restarted nature (Schoenmakers et al. 2000; Saikia & Jamroz 2009). The newly created jets in such sources travel outwards through the cocoon formed by the earlier cycle/episode of activity rather than the common intergalactic or intracluster medium, after moving through the interstellar medium of the host galaxy. In general, the outer double lobes are aligned with the inner ones and can extend from few kpc up to Mpc scales. It is likely that in DDRGs, due to an unknown mechanism or activity, the interruption of these bipolar relativistic jet flows has occurred leading to such morphologies.

In the last ~ 20 years only about 140 DDRGs have been reported (Schoenmakers et al. 2000; Nandi & Saikia 2012; Kuźmierz et al. 2017; Mahatma et al. 2019; Nandi et al. 2019). Recently, Mahatma et al. 2019 created a new sample of 33 DDRGs using the LoTSS DR1 and a follow up observations at higher frequency with the Jansky Very Large Array (JVLA), where they compared the optical and infrared magnitudes and colours of their host galaxies with a sample of normal radio galaxies. They find that the host galaxy properties of both DDRGs and normal radio galaxies are similar and suggest that

⁵ <http://www.jb.man.ac.uk/atlas/>

the DDRG activity is a regular part of the life cycle of the radio galaxies. We have found only 14 GRGs (largest sample reported till date) in our entire sample with DDRG type morphology as seen in Fig. 12 and Fig. 13, indicating that megaparsec scale DDRGs are rare. Except J110457.07+480913, J110613.55+485748.29, J121900.76+505254.41, J123754.05+512201.28, J134313.31+560008.35 and J141504.70+463428.97, rest 8 were previously classified as DDRGs in the literature (7 in Mahatma et al. 2019 and GRG J140718.48+513204.63 in Nandi & Saikia 2012).

We note that sources 28 (GRG J112130.35+494208.14) and 221 (GRG J145002.36+540528.27) show some signs of being DDRGs but based on the current radio maps available to us they cannot be confirmed as DDRGs and hence are candidate DDRGs.

These sources provide a unique opportunity to study timescales of AGN recurrent activity. As seen in Fig. 12, Fig. 13 and Fig. 14, some components of the giant DDRGs are very faint and it was only possible with LoTSS's high sensitivity and resolution to detect and sufficiently resolve them. Further studies on host AGN/galaxy properties of DDRGs along with their local environments are needed to understand these peculiar sources better.

4.3. GRGs in dense environments

It has been hypothesized that growth of GRGs to enormous size is favoured by their location in low density environments. Using our large new sample of GRGs, it is now possible to test this hypothesis for objects with low redshifts ($z < 0.55$). We find, at least 8.4% GRGs (Table 3) from our sample (20/239) are located in high density environments and are hosted by BCGs. This low number of BCG GRGs is possibly due to absence of data for high redshift clusters in the WHL cluster catalogue and GMBCG catalogue, which are more sensitive to clusters with redshift less than ~ 0.42 and ~ 0.55 , respectively (see Sect. 3.5). We have 128 GRGs with $z \leq 0.55$, therefore at least $\sim 16\%$ (20/128) GRGs are in dense cluster environments. Based on the work of Paul et al. (2017), virialized structures of mass $M_{200} \geq 0.8 \times 10^{14} M_{\odot}$ are classified as cluster of galaxies and non-virialized gravitationally bound structures consisting of few galaxies with $M_{200} < 0.8 \times 10^{14} M_{\odot}$ are classified as group of galaxies. Using this classification there are 14 GRGs in clusters and 3 GRGs in groups of galaxies (see Table 3).

Croston et al. (2019) carried out a study of environments of ~ 8000 radio loud AGNs from the LoTSS, where they find that only 10% of AGNs are associated with high density environments like galaxy groups/clusters and AGNs with $L_{150} > 10^{25} \text{ W Hz}^{-1}$ (where L_{150} is the radio power at 150 MHz) are more likely to be in cluster environments. Similarly our 20 BCG GRGs which are also radio loud AGNs in galaxy cluster, exhibit $P_{144\text{MHz}}$ or $L_{150} > 10^{25} \text{ W Hz}^{-1}$.

To compare properties of galaxy clusters which have radio loud BCGs (also BCG-GRGs) and the galaxy clusters with radio quiet BCGs, we extracted all the galaxy clusters present in the LoTSS-DR1 HETDEX region from the WHL cluster catalog. This resulted in total of 5027 galaxy clusters below redshift (z) of 0.55 (highest redshift of BCG-GRG in our sample). Next, we cross matched the location of BCGs (5027) from the WHL cluster with the VAC and FIRST catalog to determine all the radio loud BCGs in this region.

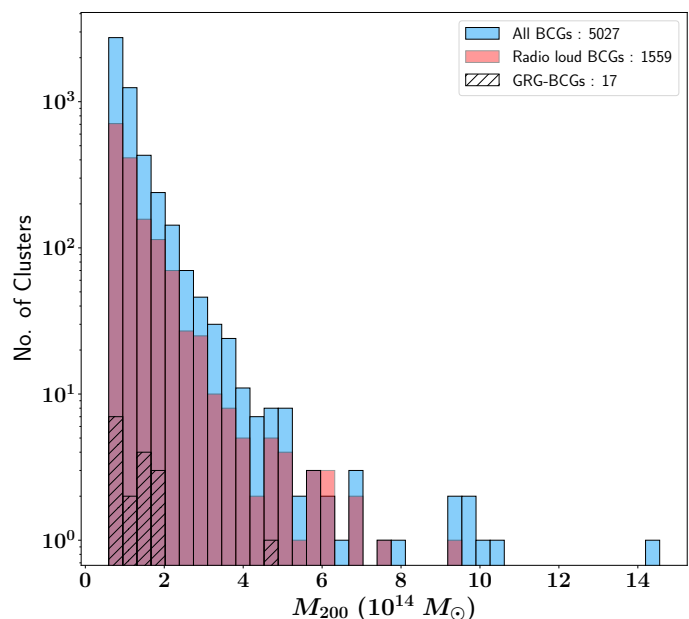


Fig. 10: Distribution of M_{200} (mass of the cluster within r_{200}) of clusters of galaxies.

The final number of galaxy clusters with radio loud BCGs is 1559 (unique sample after combining results from matches with VAC and FIRST). The 1559 radio loud BCGs also include our 17 BCG-GRGs (found in WHL; see Table 3). In Fig. 10 we have shown the distribution of M_{200} , which is the mass of the cluster within r_{200} , of clusters of galaxies, where r_{200} is the radius within which the mean density of a cluster is 200 times of the critical density of the universe. We note that about 30% (1541) clusters have radio loud BCGs which are not giants and only $\sim 0.34\%$ of the clusters in WHL cluster catalogue (HETDEX region) host a GRG (or have BCG GRG). Though the GRG sample is less here but more or less complete, it seems that GRGs which are BCGs tend to avoid very high mass clusters and prefer less dense clusters.

The BCG GRGs also can possibly trace the inhomogeneities in the intergalactic gas, which could be one of the determining factors for the growth and evolution of these giant sources. The forward propagation of jets as well as backflow from hotspots in the lobes of GRGs are influenced by the gas they encounter. Only 1/20 BCG GRGs in our sample have FR-I type radio morphology and rest 19 have FR-II type of radio morphology. Almost all these sources do not exhibit highly linear structures but are distorted or bent to certain extent, which indicates the effect of interaction between the cluster medium and the expanding radio jets of the BCG GRGs. Remarkably, despite the possible resistance presented by the cluster medium, these sources have managed to grow to megaparsec scales. This seems to suggest the presence of extraordinarily powerful AGN powering these sources.

5. Summary

A total of 239 GRGs (Table. 1) were found in $\sim 424 \text{ deg}^2$ area using LoTSS, which is just $\sim 2\%$ area of the total survey that is planned to cover the northern sky. Our sample of 239 GRGs represents a lower number estimate due to

Table 1: Short summary of classification of GRGs.

Classification	No. of objects
GRQs	40
BCGs	20
FR-II	215
FR-I	18
HyMoRS	6
DDRGs	14

limitation of optical data, which is essential for identifying host galaxy and its corresponding redshift. Assuming the isotropy and homogeneity of the Universe, if we extrapolate the number of GRGs expected to be found over the final sky coverage of LoTSS ($\sim 2\pi$ steradians), then we should be able to find at least ~ 12000 GRGs with LoTSS's sensitivity. Upcoming deep optical spectroscopic surveys like the WEAVE-LOFAR survey (Smith et al. 2016) will provide crucial data (redshifts) for identifying hundreds of more GRGs. The summary of the paper is as follows :

1. Our sample of 239 GRGs is in the redshift range of 0.1 to 2.3, out of which 225 are newly found. This makes it the largest sample discovered to date.
2. About 16% (40/239) of the sample are hosted by quasars, where three GRGs are above redshift of 2.
3. The depth and resolution of LoTSS images has enabled us to find GRGs with low powers of $\sim 10^{24}$ W Hz $^{-1}$ at 144 MHz.
4. We show that the spectral index of GRGs and GRQs is similar to that of their low sized counterparts (NRGs).
5. Majority (90%) of our sample show FR-II type of morphology.
6. We have found 14 double-double GRGs, which is $\sim 5\%$ of our sample. It is the largest reporting sample of giant DDRGs till date.
7. We have found 6 GRGs with HyMoRS morphology which are very rare.
8. Based on the optical data, we find that none of the GRGs in our sample are hosted by spiral galaxies.
9. At $z < 0.55$, at least $\sim 16\%$ of GRGs lie at the centers (BCGs) of either big galaxy groups or galaxy clusters.

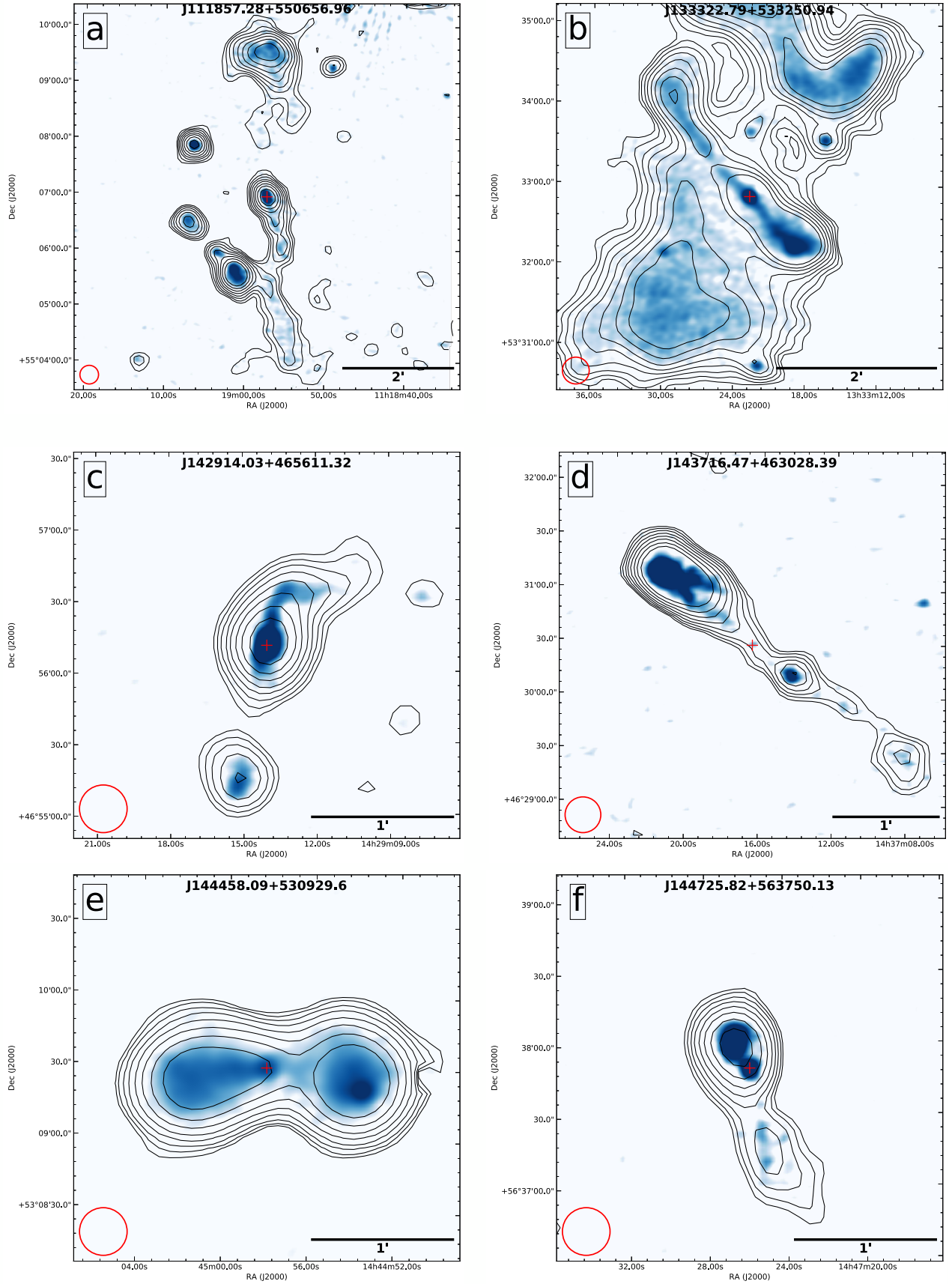


Fig. 11: LoTSS DR1 ($6'' \times 6''$) maps of 6 HyMoRS in inverted blue colour scale with ($20'' \times 20''$) black contours (eight equally spaced relative contour levels starting from three times σ , where σ denotes the local RMS noise). The circles at bottom left corner represents the beam sizes of low resolution radio maps of LoTSS DR1. The red colour marker '+' indicates the location of the host galaxy.

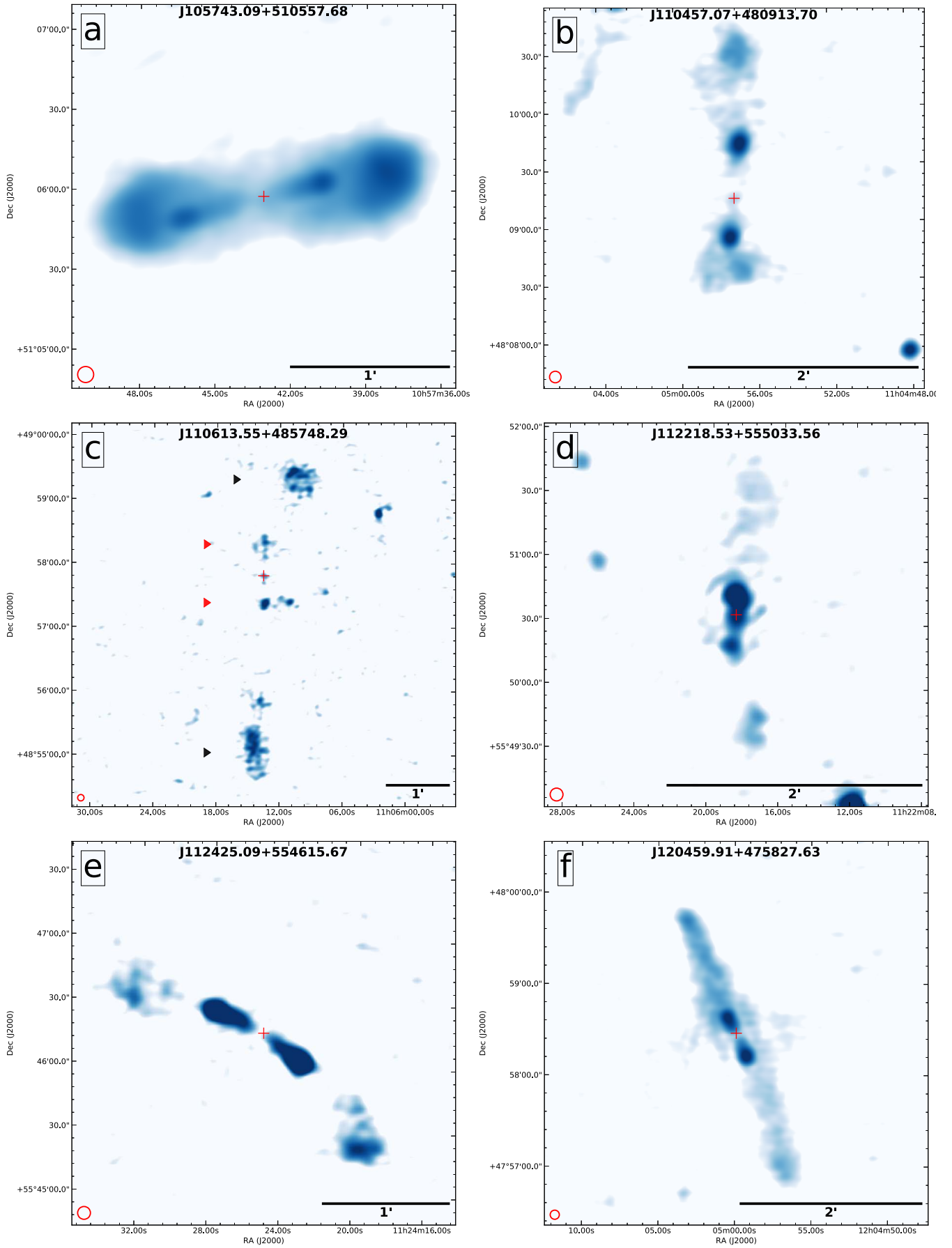


Fig. 12: DDRGs postage-1 (1-6 DDRGs) made using LoTSS DR1 144 MHz radio image with $6'' \times 6''$ resolution in inverted blue colour scale. Refer to Sect. 4.2.2 for more details. The circle at bottom left corner represents the beam size. The red colour marker '+' indicates the location of the host galaxy. In source J110613.55+485748.29 (sub-image c) the two red markers indicate the inner pair of the DDRG and the black markers indicate the out pair of the DDRG.

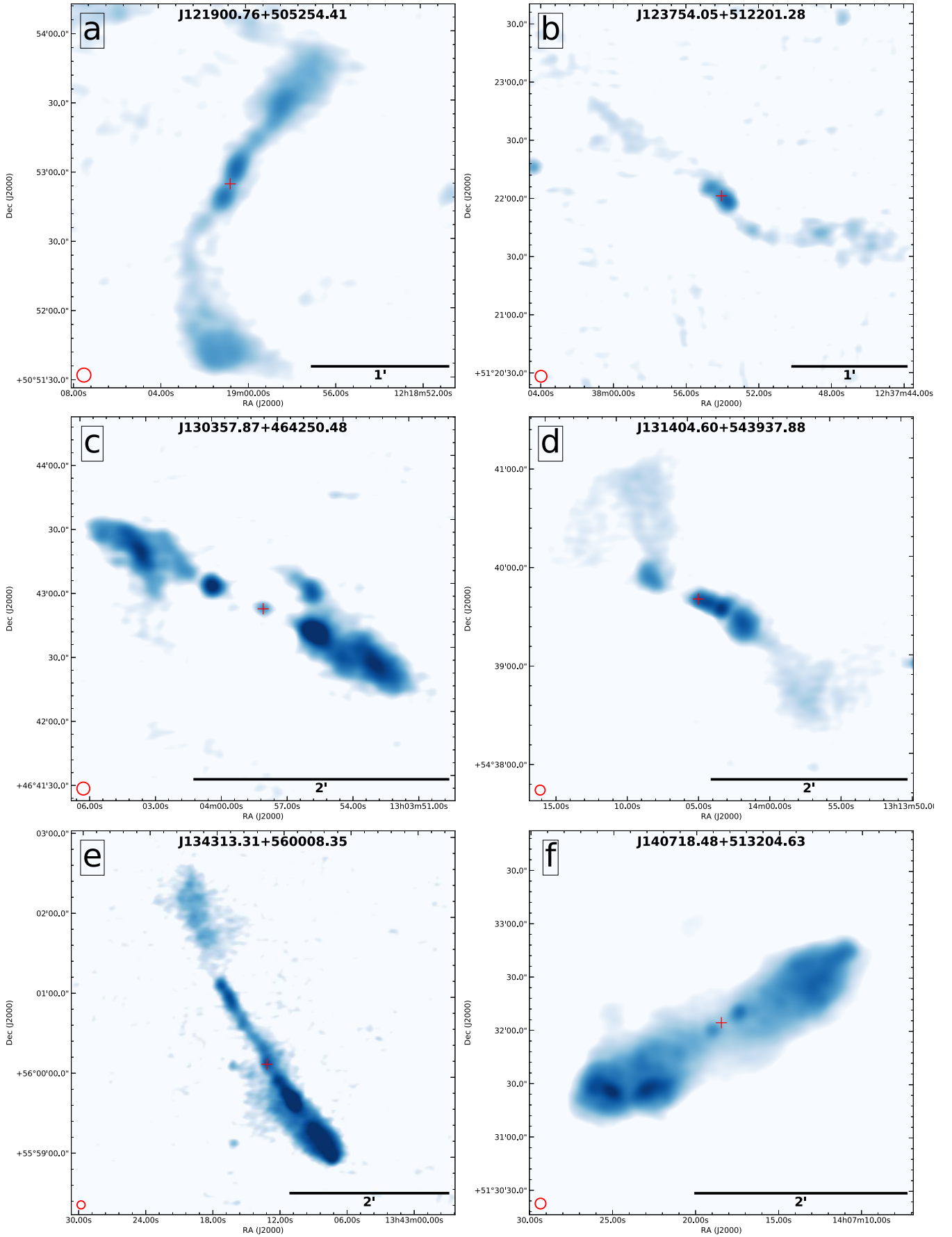


Fig. 13: DDRGs postage-1 (7-12 DDRGs) made using LoTSS DR1 144 MHz radio image with $6'' \times 6''$ resolution in inverted blue colour scale. Refer to Sect. 4.2.2 for more details.

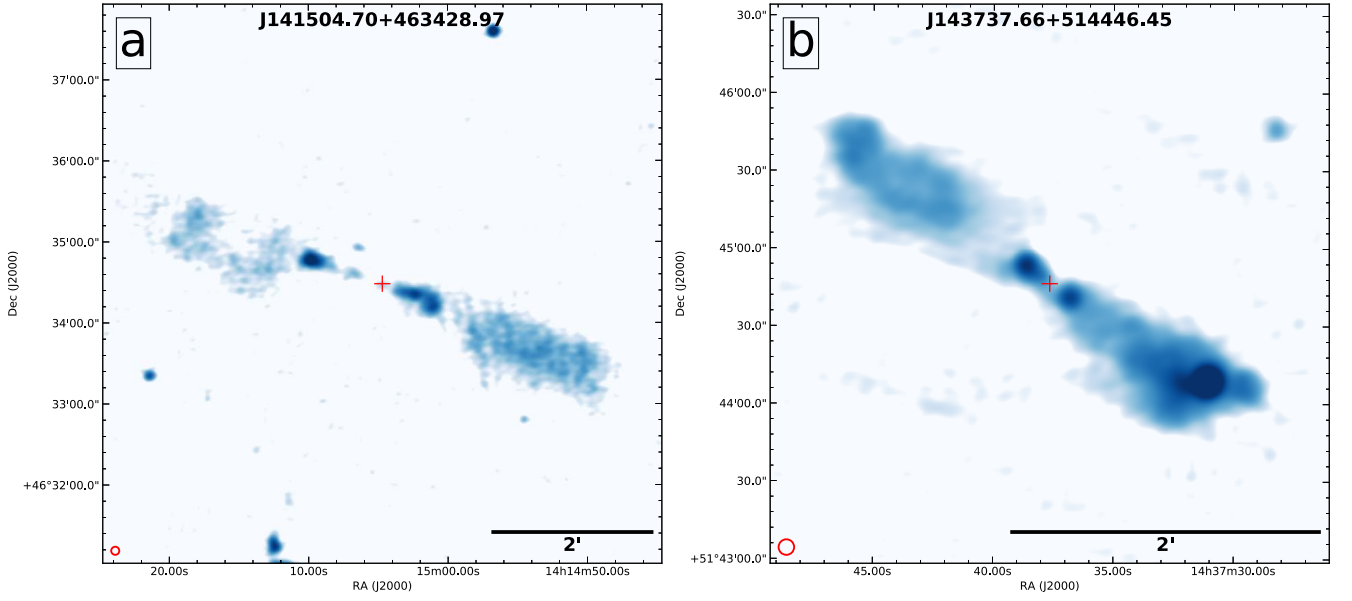


Fig. 14: DDRGs postage-1 (13-14 DDRGs) made using LoTSS DR1 144 MHz radio image with $6'' \times 6''$ resolution in inverted blue colour scale. Refer to Sect. 4.2.2 for more details.

Table 2: Here, we list all the 239 GRGs found from the LoTSS. The Columns (2)-RA & (3)-Dec are the right ascension & declination (J2000.0 epoch) which indicate the center of the host galaxies of the GRG/GRQs from either SDSS or Pan-STARRS. Column (4) ‘Class’ represents the type of host of the GRG - G:galaxy, Q:quasar and Q_c : quasar candidate. In Column (5), z (redshift) marked with \dagger represents spectroscopic redshift from SDSS, \star represents photometric redshift from SDSS and \S indicates redshifts from VAC. The one marked with superscript ‘a’ refers to redshift obtained from O’Sullivan et al. 2019 (spectroscopic) and ‘b’ from Lopes 2007 (photometric), owing to higher accuracy when compared to SDSS or VAC. Columns (6) & (7) are angular size and projected linear size of the source. Columns (8) & (9) are the integrated flux of sources at 144 MHz and its corresponding radio power. Column (10) α_{144}^{1400} is the integrated spectral index computed between 1400 MHz (NVSS) flux and 144 MHz (LoTSS) flux. Rows marked with ‘-’ in Column (10) are blended sources for which we do not present a spectral index. Column (11) indicates the morphological type of the GRG: I represents FR-I type, II represents FR-II type and III represents HyMoRS. Column (12) represents the r band magnitude of the host galaxies of the GRGs obtained from SDSS. For sources 162 and 233, in absence of reliable magnitudes from SDSS, we have given r band magnitudes from PAN-STARRS. The last column (Ref) shows references (see end of the table) for the GRGs already known in literature.

Sr.No	RA	Dec	Class	z	Size	Size	S_{144MHz}	P_{144MHz}	α_{144}^{1400}	FR	r	Ref
(1)	(HMS)	(DMS)	(4)	(5)	(")	(Mpc)	(mJy)	($10^{26} \text{ W Hz}^{-1}$)	(10)	Type	Mag	(13)
1	10 57 05.43	+53 26 24.79	Q	0.45992 \pm 0.00003 †	153	0.92	1485 \pm 297	11.00 \pm 2.20	0.69 \pm 0.09	II	19.48 \pm 0.03	-
2	10 57 09.25	+48 40 41.03	G	0.27627 \pm 0.00005 †	439	1.90	1062 \pm 212	2.55 \pm 0.51	0.79 \pm 0.09	II	18.01 \pm 0.01	-
3	10 57 14.71	+55 48 56.02	Q	1.39830 \pm 0.00166 †	91	0.79	20 \pm 4	1.84 \pm 0.40	0.64 \pm 0.11	II	18.86 \pm 0.01	-
4	10 57 21.22	+53 14 22.42	G	0.33645 \pm 0.00011 †	145	0.72	11 \pm 4	0.04 \pm 0.02	<0.63	II	18.36 \pm 0.02	-
5	10 57 25.96	+49 29 00.31	G	0.47614 \pm 0.00022 †	157	0.96	94 \pm 19	0.81 \pm 0.16	0.86 \pm 0.10	II	19.76 \pm 0.04	-
6	10 57 43.09	+51 05 57.68	G	0.46269 \pm 0.00004 †	126	0.76	657 \pm 131	5.22 \pm 1.05	0.85 \pm 0.09	II	20.59 \pm 0.05	-
7	10 58 17.90	+51 40 17.70	G	0.41497 \pm 0.00003 †	330	1.86	298 \pm 88	1.77 \pm 0.53	0.74 \pm 0.22	II	19.62 \pm 0.03	8
8	11 01 47.58	+46 49 11.20	G	0.68096 \pm 0.00017 †	150	1.09	532 \pm 106	10.70 \pm 2.15	0.88 \pm 0.09	II	21.17 \pm 0.12	3
9	11 01 59.17	+46 45 34.19	G	0.46877 \pm 0.00019 †	127	0.77	168 \pm 34	1.35 \pm 0.27	0.80 \pm 0.09	II	19.49 \pm 0.03	-
10	11 04 33.11	+46 42 25.76	G	0.14126 \pm 0.00001 †	305	0.78	253 \pm 51	0.14 \pm 0.03	0.68 \pm 0.09	I	16.68 \pm 0.01	9
11	11 04 57.07	+48 09 13.70	G	0.41471 \pm 0.00011 †	137	0.78	37 \pm 7	0.22 \pm 0.04	0.75 \pm 0.12	II	19.84 \pm 0.03	-
12	11 05 15.27	+54 41 09.30	G	0.28425 \pm 0.00010 †	169	0.75	89 \pm 18	0.23 \pm 0.05	0.90 \pm 0.10	II	20.64 \pm 0.05	-
13	11 06 13.55	+48 57 48.29	G	0.634 \pm 0.108 §	292	2.06	10 \pm 2	0.15 \pm 0.07	<0.63	II	20.25 \pm 0.12	-
14	11 08 00.77	+51 20 28.64	G	0.801 \pm 0.172 §	157	1.21	192 \pm 39	7.64 \pm 4.33	1.38 \pm 0.13	II	20.47 \pm 0.10	-
15	11 08 15.18	+50 10 03.84	G	0.66956 \pm 0.00023 †	103	0.74	29 \pm 6	0.77 \pm 0.16	<1.54	II	21.49 \pm 0.10	-
16	11 08 33.99	+48 32 03.03	Q	2.12100 \pm 0.00026 †	85	0.73	30 \pm 6	4.32 \pm 0.89	0.21 \pm 0.09	II	19.64 \pm 0.02	-
17	11 09 20.21	+48 15 00.51	G	0.39190 \pm 0.00009 †	183	1.00	12 \pm 2	0.05 \pm 0.01	0.33 \pm 0.12	II	18.31 \pm 0.02	-
18	11 09 35.40	+51 04 02.27	Q	1.18230 \pm 0.00049 †	106	0.90	186 \pm 37	12.70 \pm 2.58	-	II	19.13 \pm 0.01	2
19	11 09 36.42	+53 13 48.13	G	0.288 \pm 0.027 §	230	1.03	583 \pm 117	1.52 \pm 0.45	-	II	18.16 \pm 0.01	-
20	11 10 11.02	+53 30 58.74	G	0.58483 \pm 0.00018 †	142	0.96	565 \pm 113	7.44 \pm 1.50	-	II	20.64 \pm 0.07	-
21	11 12 59.52	+49 42 27.11	G	0.502 \pm 0.133 §	244	1.54	223 \pm 45	1.90 \pm 1.26	0.55 \pm 0.09	II	21.25 \pm 0.09	-
22	11 13 31.45	+46 22 16.04	G	0.589 \pm 0.145 §	112	0.76	340 \pm 68	4.66 \pm 2.92	0.80 \pm 0.09	II	21.84 \pm 0.10	-
23	11 14 00.10	+53 22 16.57	G	0.727 \pm 0.116 §	125	0.93	6 \pm 1	0.17 \pm 0.08	<1.16	II	21.35 \pm 0.15	-
24	11 15 29.60	+56 00 39.64	G	0.57880 \pm 0.00007 †	107	0.72	82 \pm 16	1.06 \pm 0.21	0.77 \pm 0.09	II	20.33 \pm 0.06	-
25	11 18 57.28	+55 06 56.96	G	0.35159 \pm 0.00007 †	347	1.77	669 \pm 134	2.73 \pm 0.55	-	III	19.23 \pm 0.02	-
26	11 20 12.70	+51 32 56.34	G	0.72720 \pm 0.00019 †	103	0.77	94 \pm 19	2.17 \pm 0.44	0.86 \pm 0.10	II	21.03 \pm 0.10	-
27	11 21 26.44	+53 44 56.71	G	0.10378 \pm 0.00002 †	450	0.88	399 \pm 80	0.11 \pm 0.02	0.61 \pm 0.09	II	15.11 \pm 0.00	-
28	11 21 30.35	+49 42 08.14	G	0.494 \pm 0.071 §	132	0.82	86 \pm 17	0.72 \pm 0.29	0.61 \pm 0.09	II	20.00 \pm 0.04	-
29	11 22 18.53	+55 50 33.56	G	0.90962 \pm 0.00047 †	126	1.01	56 \pm 11	2.05 \pm 0.42	-	II	21.92 \pm 0.18	-
30	11 24 25.09	+55 46 15.67	G	0.80902 \pm 0.00042 †	146	1.13	26 \pm 5	0.75 \pm 0.15	0.77 \pm 0.13	II	22.14 \pm 0.22	-
31	11 24 29.98	+46 35 23.68	Q	0.91500 \pm 0.00013 †	88	0.71	33 \pm 7	1.11 \pm 0.23	0.58 \pm 0.10	II	20.12 \pm 0.03	-
32	11 24 35.86	+49 03 25.92	G	0.47934 \pm 0.00010 †	121	0.74	120 \pm 24	1.00 \pm 0.20	-	II	20.10 \pm 0.04	-
33	11 26 29.95	+49 01 37.16	G	0.660 \pm 0.139 §	111	0.80	35 \pm 7	0.62 \pm 0.34	0.76 \pm 0.11	II	21.67 \pm 0.12	-
34	11 26 39.89	+53 34 27.20	G	0.64523 \pm 0.00024 †	125	0.89	79 \pm 16	1.30 \pm 0.26	-	II	20.49 \pm 0.08	-
35	11 27 13.18	+51 13 26.35	G	0.36132 \pm 0.00007 †	202	1.05	54 \pm 11	0.23 \pm 0.05	0.60 \pm 0.10	I	18.12 \pm 0.01	-
36	11 27 42.06	+52 19 44.68	G	0.72484 \pm 0.00031 †	114	0.85	6 \pm 1	0.13 \pm 0.03	<0.82	II	21.75 \pm 0.12	-
37	11 28 54.65	+56 20 09.50	G	0.594 \pm 0.067 §	188	1.29	115 \pm 23	1.57 \pm 0.53	-	II	21.12 \pm 0.10	-
38	11 29 43.41	+51 24 12.11	G	0.79923 \pm 0.00008 †	226	1.75	70 \pm 14	2.15 \pm 0.43	0.94 \pm 0.11	II	21.53 \pm 0.15	-
39	11 32 02.31	+47 28 24.14	G	0.26431 \pm 0.00006 †	205	0.86	38 \pm 8	0.08 \pm 0.02	0.61 \pm 0.10	II	17.82 \pm 0.01	-
40	11 32 50.67	+50 57 04.68	G	0.35857 \pm 0.00010 †	138	0.71	63 \pm 13	0.26 \pm 0.05	0.65 \pm 0.10	II	18.53 \pm 0.01	-
41	11 34 21.89	+56 34 13.28	G	0.626 \pm 0.092 §	227	1.59	21 \pm 4	0.35 \pm 0.14	0.92 \pm 0.16	II	20.98 \pm 0.08	-
42	11 34 35.48	+46 08 00.45	G	0.704 \pm 0.163 §	300	2.21	41 \pm 8	0.90 \pm 0.54	0.87 \pm 0.16	II	21.39 \pm 0.08	-
43	11 35 03.20	+48 26 12.12	G	0.22597 \pm 0.00006 †	202	0.76	131 \pm 26	0.20 \pm 0.04	0.78 \pm 0.09	II	17.59 \pm 0.01	-
44	11 39 31.77	+47 21 24.3	G	0.51791 \pm 0.00012 †	312	2.00	126 \pm 25	1.23 \pm 0.25	0.70 \pm 0.09	II	20.32 \pm 0.05	-
45	11 43 05.55	+52 27 26.89	Q	1.67249 \pm 0.00082 †	86	0.75	100 \pm 20	14.90 \pm 3.04	0.72 \pm 0.09	II	20.23 \pm 0.03	-
46	11 48 14.98	+54 57 16.49	G	0.22679 \pm 0.00003 †	202	0.76	22 \pm 4	0.03 \pm 0.01	0.51 \pm 0.11	I	17.26 \pm 0.01	-
47	11 51 59.9	+49 50 56.11	Q	0.89148 \pm 0.00017 †	96	0.77	1015 \pm 203	36.60 \pm 7.37	0.78 \pm 0.09	II	18.38 \pm 0.01	-
48	11 52 16.89	+46 24 51.37	G	0.44454 \pm 0.02004 *	148	0.87	43 \pm 9	0.28 \pm 0.06	0.56 \pm 0.10	II	19.68 \pm 0.03	-
49	11 58 39.12	+46 55 45.35	G	0.65974 \pm 0.00023 †	128	0.92	18 \pm 4	0.36 \pm 0.07	<1.08	II	21.61 \pm 0.13	-
50	12 01 18.14	+46 32 41.98	G	0.638 \pm 0.088 §	129	0.91	111 \pm 22	1.87 \pm 0.73	0.83 \pm 0.10	II	21.82 \pm 0.15	-

Table 2: continued.

Sr.No	RA (HMS)	Dec (DMS)	Class	z	Size ($''$)	Size (Mpc)	S_{144MHz} (mJy)	P_{144MHz} (10^{26}W Hz^{-1})	α_{144}^{1400}	FR Type	r Mag	Ref
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)
51	12 01 22.84	+49 25 37.14	G	0.205 \pm 0.045 [§]	345	1.20	1946 \pm 389	2.56 \pm 1.35	1.09 \pm 0.09	I	18.14 \pm 0.01	-
52	12 02 06.86	+51 41 34.63	G	0.61278 \pm 0.00028 [†]	141	0.98	87 \pm 17	1.27 \pm 0.26	-	II	21.11 \pm 0.10	-
53	12 04 59.91	+47 58 27.63	G	0.52963 \pm 0.08487 [*]	192	1.24	138 \pm 28	1.53 \pm 0.66	0.89 \pm 0.09	II	20.96 \pm 0.07	-
54	12 05 28.50	+56 13 42.32	G	0.681 \pm 0.121 [§]	118	0.86	26 \pm 5	0.48 \pm 0.23	<0.74	II	20.54 \pm 0.13	-
55	12 05 46.98	+50 48 59.94	G	0.69799 \pm 0.00021 [†]	160	1.18	352 \pm 70	7.60 \pm 1.53	0.90 \pm 0.09	II	21.02 \pm 0.09	-
56	12 08 26.14	+47 15 25.03	G	0.56407 \pm 0.00021 [†]	113	0.75	34 \pm 7	0.41 \pm 0.08	-	II	20.76 \pm 0.06	-
57	12 10 46.06	+53 29 22.90	G	0.44842 \pm 0.00010 [†]	119	0.71	234 \pm 47	1.73 \pm 0.35	0.85 \pm 0.09	I	19.87 \pm 0.03	-
58	12 10 55.67	+49 49 39.77	G	0.53442 \pm 0.05992 [*]	110	0.72	49 \pm 10	0.53 \pm 0.18	0.79 \pm 0.10	I	20.57 \pm 0.06	-
59	12 11 40.91	+50 10 15.11	G	0.33985 \pm 0.14011 [*]	143	0.71	372 \pm 74	1.45 \pm 1.42	0.86 \pm 0.09	II	21.78 \pm 0.13	-
60	12 13 04.34	+51 18 26.38	Q	0.61680 \pm 0.00013 [†]	118	0.82	102 \pm 20	1.49 \pm 0.30	0.72 \pm 0.09	II	19.78 \pm 0.02	-
61	12 15 55.53	+51 24 16.41	G	0.44477 \pm 0.00038 [†]	277	1.63	74 \pm 15	0.52 \pm 0.10	-	I	19.17 \pm 0.02	-
62	12 17 06.80	+51 47 34.33	G	0.616 \pm 0.089 [§]	183	1.28	57 \pm 11	0.95 \pm 0.39	0.99 \pm 0.12	I	21.61 \pm 0.10	-
63	12 17 07.40	+48 58 22.48	G	0.484 \pm 0.029 [§]	136	0.84	172 \pm 34	1.46 \pm 0.36	-	I	19.60 \pm 0.03	-
64	12 18 18.15	+53 27 21.33	G	0.56754 \pm 0.04447 [*]	183	1.23	1270 \pm 254	16.00 \pm 4.41	0.81 \pm 0.09	II	21.77 \pm 0.17	-
65	12 18 49.88	+50 26 17.59	G	0.19920 \pm 0.00002 [†]	210	0.71	2942 \pm 588	3.34 \pm 0.67	0.67 \pm 0.09	II	17.12 \pm 0.01	6
66	12 19 00.76	+50 52 54.41	G	0.38509 \pm 0.00008 [†]	142	0.77	78 \pm 16	0.38 \pm 0.08	0.68 \pm 0.10	I	18.35 \pm 0.01	-
67	12 19 35.92	+46 59 29.88	G	0.67865 \pm 0.00032 [†]	114	0.83	19 \pm 4	0.35 \pm 0.07	-	II	21.14 \pm 0.09	-
68	12 19 52.32	+47 20 58.49	Q	0.65310 \pm 0.00009 [†]	241	1.72	115 \pm 23	2.06 \pm 0.41	0.85 \pm 0.10	II	19.10 \pm 0.01	-
69	12 19 52.55	+46 54 25.74	Q	1.85116 \pm 0.00041 [†]	88	0.76	60 \pm 12	8.28 \pm 1.69	0.43 \pm 0.09	II	18.62 \pm 0.01	-
70	12 20 07.85	+47 05 19.65	G	0.689 \pm 0.159 [§]	149	1.09	109 \pm 22	2.93 \pm 1.75	1.37 \pm 0.17	II	21.97 \pm 0.12	-
71	12 20 28.13	+52 51 44.89	G	0.34583 \pm 0.00008 [†]	255	1.29	58 \pm 12	0.21 \pm 0.04	0.55 \pm 0.09	II	18.09 \pm 0.01	-
72	12 22 55.24	+49 26 42.32	G	0.20315 \pm 0.00004 [†]	354	1.22	84 \pm 17	0.10 \pm 0.02	0.48 \pm 0.09	II	17.21 \pm 0.01	-
73	12 23 27.25	+52 36 45.38	G	0.62112 \pm 0.05830 [*]	111	0.77	59 \pm 12	0.89 \pm 0.27	0.74 \pm 0.10	I	21.19 \pm 0.13	-
74	12 25 03.66	+47 23 36.40	G	0.774 \pm 0.154 [§]	95	0.73	14 \pm 3	0.32 \pm 0.17	0.51 \pm 0.12	II	21.26 \pm 0.09	-
75	12 25 31.36	+49 46 43.95	G	0.32488 \pm 0.00009 [†]	206	1.00	21 \pm 4	0.07 \pm 0.01	0.60 \pm 0.12	I	18.32 \pm 0.01	-
76	12 25 44.39	+49 49 42.93	G	0.63224 \pm 0.08754 [*]	164	1.16	286 \pm 57	5.11 \pm 2.00	1.00 \pm 0.09	II	21.06 \pm 0.06	-
77	12 25 58.20	+53 09 38.38	G	0.81100 \pm 0.07937 [*]	147	1.14	263 \pm 53	8.15 \pm 2.55	0.91 \pm 0.09	II	20.25 \pm 0.15	-
78	12 28 26.35	+52 31 01.10	G	0.453 \pm 0.106 [§]	157	0.93	106 \pm 21	0.78 \pm 0.46	0.80 \pm 0.09	II	20.65 \pm 0.06	-
79	12 29 36.25	+50 13 04.65	G	0.38357 \pm 0.00007 [†]	198	1.07	56 \pm 11	0.27 \pm 0.05	0.61 \pm 0.10	I	18.53 \pm 0.02	-
80	12 29 59.59	+53 32 47.04	G	0.52300 \pm 0.05280 [*]	137	0.88	296 \pm 59	3.20 \pm 1.01	0.90 \pm 0.09	II	21.12 \pm 0.11	-
81	12 30 14.05	+54 11 41.14	G	0.610 \pm 0.087 [§]	205	1.42	7 \pm 1	0.09 \pm 0.04	<0.57	II	20.96 \pm 0.10	-
82	12 32 04.95	+53 06 27.31	G	0.20604 \pm 0.00006 [†]	274	0.95	120 \pm 24	0.15 \pm 0.03	0.90 \pm 0.10	II	17.31 \pm 0.01	-
83	12 32 50.45	+49 06 26.14	G	0.69015 \pm 0.00013 [†]	256	1.87	205 \pm 41	3.99 \pm 0.80	0.76 \pm 0.09	II	20.92 \pm 0.06	-
84	12 33 05.45	+49 02 51.93	Q	1.35200 \pm 0.00107 [†]	104	0.89	332 \pm 66	31.00 \pm 6.30	-	II	20.55 \pm 0.03	-
85	12 35 01.52	+53 17 55.09	G	0.34480 \pm 0.00030 [†]	683	3.44	859 \pm 172	3.46 \pm 0.69	0.85 \pm 0.09	II	19.24 \pm 0.04	-
86	12 35 48.77	+50 50 36.30	G	0.650 \pm 0.131 [§]	105	0.75	84 \pm 17	1.41 \pm 0.75	-	II	21.49 \pm 0.10	-
87	12 36 48.37	+46 04 05.91	G	0.615 \pm 0.154 [§]	125	0.87	268 \pm 54	3.96 \pm 2.53	0.75 \pm 0.09	II	21.39 \pm 0.10	-
88	12 37 54.05	+51 22 01.28	G	0.39184 \pm 0.00018 [†]	164	0.89	25 \pm 5	0.13 \pm 0.03	-	II	19.68 \pm 0.03	-
89	12 38 07.78	+53 25 55.91	Q	0.34680 \pm 0.00004 [†]	157	0.80	455 \pm 91	1.73 \pm 0.35	0.61 \pm 0.09	II	17.29 \pm 0.01	-
90	12 39 33.21	+50 07 08.01	Q	2.31500 \pm 0.00024 [†]	89	0.75	3 \pm 1	2.32 \pm 0.52	<1.39	II	18.49 \pm 0.01	-
91	12 40 12.46	+53 34 37.25	Q	0.29300 \pm 0.00003 [†]	164	0.74	1527 \pm 305	4.15 \pm 0.83	0.76 \pm 0.09	II	18.52 \pm 0.01	1
92	12 40 31.72	+48 59 36.89	G	0.53607 \pm 0.04030 [*]	136	0.89	605 \pm 121	6.35 \pm 1.72	0.70 \pm 0.09	II	21.17 \pm 0.08	-
93	12 41 42.34	+51 35 14.32	G	0.52951 \pm 0.00011 [†]	161	1.04	177 \pm 36	1.86 \pm 0.37	0.76 \pm 0.09	II	19.57 \pm 0.04	-
94	12 45 25.16	+47 07 10.11	G	0.51179 \pm 0.00015 [†]	111	0.70	22 \pm 5	0.20 \pm 0.04	0.60 \pm 0.10	II	19.95 \pm 0.03	-
95	12 45 31.06	+48 50 20.16	G	0.25359 \pm 0.00010 [†]	174	0.71	40 \pm 8	0.08 \pm 0.02	0.69 \pm 0.11	I	17.99 \pm 0.01	-
96	12 49 03.46	+52 40 05.29	G	0.663 \pm 0.107 [§]	135	0.97	66 \pm 13	1.27 \pm 0.56	0.93 \pm 0.10	II	21.60 \pm 0.13	-
97	12 49 13.20	+50 00 43.68	G	0.34956 \pm 0.00010 [†]	150	0.76	295 \pm 59	1.13 \pm 0.23	0.59 \pm 0.09	I	18.82 \pm 0.02	-
98	12 51 42.04	+50 34 24.65	G	0.54904 \pm 0.00007 [†]	150	0.99	8355 \pm 1671	99.90 \pm 20.10	0.86 \pm 0.09	II	19.35 \pm 0.03	5
99	12 56 24.74	+55 28 25.99	G	0.808 \pm 0.179 [§]	113	0.87	1765 \pm 353	55.50 \pm 32.10	0.94 \pm 0.09	II	21.01 \pm 0.12	-
100	12 57 16.35	+51 17 58.11	Q	0.52582 \pm 0.00004 [†]	170	1.09	96 \pm 19	0.98 \pm 0.20	0.74 \pm 0.10	II	17.69 \pm 0.01	-
101	12 58 21.37	+54 20 29.35	G	0.612 \pm 0.099 [§]	143	0.99	521 \pm 104	7.64 \pm 3.36	-	II	21.11 \pm 0.09	-
102	12 59 50.83	+53 07 09.03	G	0.60663 \pm 0.04364 [*]	105	0.73	67 \pm 13	1.06 \pm 0.28	0.94 \pm 0.11	II	20.93 \pm 0.09	-
103	13 01 34.99	+54 08 09.21	G	0.313 \pm 0.081 [§]	168	0.79	1581 \pm 316	5.00 \pm 3.15	0.77 \pm 0.09	II	19.23 \pm 0.02	1
104	13 03 31.08	+53 59 48.65	G	0.24201 \pm 0.00005 [†]	248	0.97	50 \pm 10	0.09 \pm 0.02	0.84 \pm 0.13	II	17.24 \pm 0.01	-
105	13 03 32.19	+52 20 01.98	G	0.27200 \pm 0.00007 [†]	165	0.71	83 \pm 17	0.18 \pm 0.04	0.56 \pm 0.09	I	17.98 \pm 0.01	-
106	13 03 57.87	+46 42 50.48	G	0.58430 \pm 0.00021 [†]	159	1.08	137 \pm 27	1.87 \pm 0.38	0.83 \pm 0.09	II	20.38 \pm 0.06	-
107	13 04 30.80	+50 41 08.02	Q	0.93000 \pm 0.00039 [†]	122	0.98	119 \pm 24	4.53 \pm 0.92	0.71 \pm 0.09	II	20.33 \pm 0.03	-
108	13 05 21.32	+49 51 42.36	Q	1.25098 \pm 0.00026 [†]	88	0.75	544 \pm 109	44.80 \pm 9.07	0.82 \pm 0.09	II	18.74 \pm 0.01	-
109	13 06 35.08	+55 36 44.43	G	0.461 \pm 0.031 [§]	136	0.82	47 \pm 9	0.36 \pm 0.09	0.74 \pm 0.11	II	20.17 \pm 0.03	-
110	13 07 48.36	+56 11 46.96	G	0.83043 \pm 0.08339 [*]	120	0.94	143 \pm 29	4.28 \pm 1.37	-	II	22.23 \pm 0.29	-
111	13 09 25.74	+53 48 17.26	G	0.81187 \pm 0.00032 [†]	118	0.92	75 \pm 15	2.31 \pm 0.47	0.89 \pm 0.10	II	22.20 \pm 0.26	-
112	13 10 28.87	+52 13 40.43	G	0.650 \pm 0.094 [§]	197	1.41	451 \pm 90	7.45 \pm 3.01	0.71 \pm 0.09	II	20.28 \pm 0.06	-
113	13 12 11.14	+48 09 25.26	Q	0.71510 \pm 0.00003 [†]	110	0.82	559 \pm 112	10.30 \pm 2.08	0.51 \pm 0.09	II	17.13 \pm 0.00	-
114	13 12 16.28	+48 47 45.40	G	0.36430 \pm 0.00007 [†]	210	1.09	122 \pm 24	0.52 \pm 0.10	0.60 \pm 0.09	II	18.64 \pm 0.01	-

Table 2: continued.

Sr.No	RA (HMS)	Dec (DMS)	Class	z	Size ($''$)	Size (Mpc)	S_{144MHz} (mJy)	P_{144MHz} (10^{26}W Hz^{-1})	α_{144}^{1400}	FR Type	r Mag	Ref
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)
119	13 19 05.75	+47 46 39.23	G	0.90300 \pm 0.00010 †	262	2.10	44 \pm 9	1.70 \pm 0.34	0.83 \pm 0.12	II	20.94 \pm 0.04	-
120	13 20 29.67	+49 16 47.11	Q	0.68447 \pm 0.00009 †	113	0.82	74 \pm 15	1.40 \pm 0.28	-	II	19.08 \pm 0.01	-
121	13 22 29.07	+50 48 44.53	G	0.18439 \pm 0.00004 †	326	1.04	94 \pm 19	0.09 \pm 0.02	-	II	17.23 \pm 0.01	-
122	13 23 36.36	+47 29 49.68	G	0.440 \pm 0.035 §	180	1.05	249 \pm 50	1.69 \pm 0.46	-	II	19.36 \pm 0.03	-
123	13 23 54.44	+46 31 42.89	G	0.35737 \pm 0.00022 †	136	0.70	118 \pm 24	0.49 \pm 0.10	0.69 \pm 0.09	II	19.57 \pm 0.03	-
124	13 24 35.19	+50 41 02.31	G	0.28711 \pm 0.00006 †	162	0.72	495 \pm 99	1.41 \pm 0.28	1.14 \pm 0.09	II	18.17 \pm 0.01	-
125	13 25 54.31	+55 19 36.23	G	1.692 \pm 0.089 §	87	0.75	31 \pm 6	5.40 \pm 1.32	0.85 \pm 0.13	II	20.48 \pm 0.03	-
126	13 26 14.35	+49 34 31.52	G	0.340 \pm 0.044 §	169	0.85	3885 \pm 777	14.70 \pm 5.35	0.76 \pm 0.09	II	19.17 \pm 0.02	7
127	13 26 15.35	+48 04 22.67	Q	0.59405 \pm 0.00007 †	103	0.71	287 \pm 57	4.12 \pm 0.83	0.85 \pm 0.09	II	20.90 \pm 0.08	-
128	13 26 46.56	+53 58 17.38	Q	0.40975 \pm 0.00005 †	178	1.00	444 \pm 89	2.65 \pm 0.53	0.84 \pm 0.09	II	19.56 \pm 0.03	-
129	13 27 13.33	+52 26 49.93	G	0.53770 \pm 0.00015 †	109	0.71	10 \pm 3	0.11 \pm 0.03	<0.80	II	20.31 \pm 0.04	-
130	13 29 14.05	+52 09 38.80	G	0.620 \pm 0.070 §	118	0.82	25 \pm 5	0.40 \pm 0.14	0.86 \pm 0.16	II	20.92 \pm 0.09	-
131	13 29 24.63	+49 00 15.32	G	0.286 \pm 0.081 §	253	1.12	67 \pm 13	0.17 \pm 0.12	0.74 \pm 0.11	II	19.08 \pm 0.02	-
132	13 30 41.81	+48 27 54.88	G	0.33166 \pm 0.00010 †	153	0.75	47 \pm 10	0.16 \pm 0.03	0.53 \pm 0.10	II	18.38 \pm 0.01	-
133	13 31 35.25	+45 59 55.53	G	0.38482 \pm 0.00009 †	131	0.71	115 \pm 23	0.57 \pm 0.11	0.74 \pm 0.09	II	18.06 \pm 0.01	-
134	13 32 53.76	+48 45 32.05	G	0.73036 \pm 0.00020 †	122	0.91	65 \pm 13	1.34 \pm 0.27	0.61 \pm 0.09	II	21.27 \pm 0.11	-
135	13 32 58.28	+53 53 55.60	G	0.32261 \pm 0.00009 †	174	0.84	199 \pm 40	0.70 \pm 0.14	0.94 \pm 0.09	II	18.58 \pm 0.01	-
136	13 33 22.79	+53 32 50.94	G	0.354 \pm 0.034 §	173	0.88	423 \pm 86	1.75 \pm 0.54	-	III	18.90 \pm 0.02	-
137	13 34 11.70	+55 01 24.87	Q	1.24470 \pm 0.00043 †	91	0.78	3380 \pm 676	294.00 \pm 59.50	0.90 \pm 0.09	II	17.97 \pm 0.01	4
138	13 34 18.63	+48 13 17.08	Q	2.20849 \pm 0.00009 †	88	0.75	487 \pm 97	140.00 \pm 28.90	-	II	18.15 \pm 0.01	-
139	13 35 12.01	+49 44 27.20	G	0.436 \pm 0.041 §	173	1.01	13 \pm 3	0.08 \pm 0.03	<0.71	II	19.32 \pm 0.02	-
140	13 36 12.67	+54 47 42.24	Q	0.71180 \pm 0.00012 †	142	1.05	95 \pm 19	1.97 \pm 0.40	-	II	17.72 \pm 0.01	-
141	13 36 18.75	+53 39 52.12	G	0.30138 \pm 0.00006 †	305	1.41	49 \pm 10	0.14 \pm 0.03	0.72 \pm 0.11	II	17.91 \pm 0.01	-
142	13 36 37.92	+55 40 33.27	G	0.86400 \pm 0.00003 †	105	0.83	291 \pm 58	9.85 \pm 1.98	0.80 \pm 0.09	II	21.16 \pm 0.05	-
143	13 38 04.28	+46 46 41.19	Q	1.36900 \pm 0.00030 †	140	1.21	95 \pm 19	9.17 \pm 1.87	-	II	21.13 \pm 0.06	-
144	13 39 22.70	+50 57 47.40	G	0.316 \pm 0.098 §	174	0.83	56 \pm 11	0.17 \pm 0.13	0.66 \pm 0.10	II	19.89 \pm 0.02	-
145	13 41 03.08	+49 15 59.99	G	0.74672 \pm 0.08486 *	113	0.85	334 \pm 67	7.64 \pm 2.62	0.71 \pm 0.09	II	21.63 \pm 0.10	-
146	13 42 06.98	+47 25 53.04	G	0.17213 \pm 0.00004 †	369	1.11	366 \pm 73	0.30 \pm 0.06	0.68 \pm 0.09	II	15.92 \pm 0.00	-
147	13 43 13.31	+56 00 08.35	G	0.48474 \pm 0.00007 †	245	1.51	174 \pm 35	1.62 \pm 0.33	0.97 \pm 0.10	II	20.00 \pm 0.03	-
148	13 44 15.65	+48 45 48.96	G	0.725 \pm 0.160 §	202	1.51	495 \pm 99	11.50 \pm 6.59	0.87 \pm 0.09	II	21.66 \pm 0.14	-
149	13 44 41.82	+50 22 57.29	G	0.76317 \pm 0.00025 †	123	0.93	61 \pm 12	1.58 \pm 0.32	0.86 \pm 0.10	II	21.82 \pm 0.16	-
150	13 44 52.71	+48 57 49.96	G	0.34021 \pm 0.00011 †	167	0.84	36 \pm 7	0.14 \pm 0.03	0.92 \pm 0.13	II	19.18 \pm 0.02	-
151	13 45 57.55	+54 03 16.62	G	0.16250 \pm 0.00003 †	334	0.96	2582 \pm 516	1.95 \pm 0.39	0.89 \pm 0.09	II	17.15 \pm 0.01	2
152	13 46 00.30	+53 18 41.21	G	0.536 \pm 0.096 §	156	1.01	10 \pm 2	0.12 \pm 0.06	<1.09	II	20.61 \pm 0.07	-
153	13 46 13.49	+50 35 07.24	G	0.69605 \pm 0.00034 †	141	1.04	12 \pm 3	0.22 \pm 0.05	<0.59	II	21.49 \pm 0.12	-
154	13 46 45.96	+47 27 19.03	G	0.420 \pm 0.049 §	184	1.05	148 \pm 30	0.91 \pm 0.31	-	II	19.83 \pm 0.03	-
155	13 47 30.56	+47 05 09.28	G	0.371 \pm 0.047 §	211	1.11	77 \pm 15	0.38 \pm 0.14	1.00 \pm 0.12	II	19.53 \pm 0.02	-
156	13 48 16.94	+49 50 24.80	G	0.64318 \pm 0.00017 †	120	0.85	8 \pm 2	0.13 \pm 0.03	<0.68	II	20.24 \pm 0.06	-
157	13 48 37.68	+47 08 00.01	Q	0.50724 \pm 0.00009 †	119	0.75	127 \pm 25	1.15 \pm 0.23	0.65 \pm 0.09	II	19.84 \pm 0.05	-
158	13 48 53.20	+46 45 50.63	Q	1.66716 \pm 0.00065 †	102	0.89	58 \pm 12	11.40 \pm 2.34	1.01 \pm 0.11	II	20.37 \pm 0.03	-
159	13 49 27.92	+46 20 15.11	G	0.42110 \pm 0.00010 †	131	0.75	120 \pm 24	0.74 \pm 0.15	0.76 \pm 0.09	II	19.05 \pm 0.02	-
160	13 52 53.13	+46 25 20.66	G	0.471 \pm 0.071 §	220	1.34	397 \pm 79	3.13 \pm 1.29	0.72 \pm 0.09	II	20.63 \pm 0.04	-
161	13 54 14.72	+49 13 15.2	G	0.43885 \pm 0.13902 *	123	0.72	633 \pm 127	4.55 \pm 3.53	0.91 \pm 0.09	II	22.08 \pm 0.18	-
162	13 56 28.50	+52 42 19.23	G	0.42709 \pm 0.00011 †	297	1.71	64 \pm 13	0.37 \pm 0.07	0.50 \pm 0.09	II	18.62 \pm 0.02	-
163	13 56 35.89	+56 09 44.77	G	0.66300 \pm 0.00021 †	111	0.80	18 \pm 4	0.32 \pm 0.06	0.76 \pm 0.13	II	20.83 \pm 0.07	-
164	13 59 51.16	+47 03 21.03	G	0.46224 \pm 0.00010 †	140	0.85	23 \pm 5	0.17 \pm 0.03	0.55 \pm 0.14	II	19.48 \pm 0.03	-
165	14 02 55.41	+51 27 28.60	G	0.51796 \pm 0.16062 *	135	0.87	304 \pm 61	3.07 \pm 2.36	0.79 \pm 0.09	II	20.93 \pm 0.06	-
166	14 03 15.11	+51 44 44.77	G	0.48517 \pm 0.16607 *	228	1.41	1596 \pm 319	14.30 \pm 12.00	0.87 \pm 0.09	II	21.07 \pm 0.06	-
167	14 04 08.61	+46 42 39.42	G	0.537 \pm 0.060 §	453	2.96	224 \pm 45	2.43 \pm 0.81	-	I	20.34 \pm 0.06	-
168	14 05 40.97	+54 10 55.42	G	0.76118 \pm 0.13989 *	116	0.88	588 \pm 118	14.10 \pm 6.96	0.72 \pm 0.09	II	20.70 \pm 0.04	-
169	14 06 05.42	+55 47 49.87	G	0.78025 \pm 0.06341 *	114	0.87	598 \pm 120	16.50 \pm 4.67	0.86 \pm 0.09	II	23.09 \pm 0.40	-
170	14 06 12.30	+53 38 34.64	G	0.46455 \pm 0.00010 †	117	0.71	5 \pm 2	0.04 \pm 0.02	<0.55	II	19.90 \pm 0.03	-
171	14 07 18.48	+51 32 04.63	G	0.34048 \pm 0.00003 †	177	0.89	4246 \pm 849	16.50 \pm 3.31	0.83 \pm 0.09	II	18.52 \pm 0.02	-
172	14 07 19.96	+55 06 01.29	G	0.33224 \pm 0.00010 †	180	0.88	776 \pm 155	2.79 \pm 0.56	-	II	17.09 \pm 0.02	-
173	14 08 32.49	+47 38 37.42	Q	1.43633 \pm 0.00052 †	85	0.74	282 \pm 56	32.50 \pm 6.60	0.83 \pm 0.09	II	18.90 \pm 0.01	-
174	14 09 36.64	+53 40 27.64	G	0.52484 \pm 0.00015 †	240	1.54	819 \pm 164	9.06 \pm 1.82	0.93 \pm 0.09	II	19.98 \pm 0.05	-
175	14 11 17.03	+47 10 29.47	G	0.39350 \pm 0.00013 †	300	1.64	62 \pm 12	0.30 \pm 0.06	0.53 \pm 0.10	I	18.78 \pm 0.02	-
176	14 11 51.97	+55 09 48.74	Q	1.89670 \pm 0.00019 †	99	0.86	559 \pm 112	126.00 \pm 25.80	0.85 \pm 0.09	II	18.83 \pm 0.01	-
177	14 14 08.45	+48 41 56.11	G	1.361 \pm 0.130 §	107	0.92	1055 \pm 211	93.40 \pm 29.00	0.67 \pm 0.09	II	17.90 \pm 0.01	-
178	14 15 04.70	+46 34 28.97	G	0.220 \pm 0.029 §	332	1.22	129 \pm 26	0.18 \pm 0.07	0.65 \pm 0.09	II	18.13 \pm 0.01	-
179	14 15 54.37	+49 09 21.15	Q	1.37310 \pm 0.00037 †	111	0.96	141 \pm 28	14.40 \pm 2.93	0.82 \pm 0.09	II	19.03 \pm 0.01	2
180	14 16 06.64	+47 09 27.72	G	0.534 \pm 0.131 §	117	0.76	3 \pm 1	0.03 \pm 0.02	<0.51	II	20.30 \pm 0.07	-
181	14 16 25.89	+54 25 45.85	G	0.24474 \pm 0.00005 †	336	1.33	797 \pm 159	1.59 \pm 0.32	1.20 \pm 0.09	II	18.02 <	

Table 2: continued.

Sr.No	RA (HMS)	Dec (DMS)	Class	z	Size ($''$)	Size (Mpc)	S_{144MHz} (mJy)	P_{144MHz} (10^{26}W Hz^{-1})	α_{144}^{1400}	FR Type	r Mag	Ref
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)
187	14 23 14.59	+47 51 26.99	G	$0.666 \pm 0.110^{\S}$	197	1.42	122 ± 24	2.18 ± 0.98	-	II	20.74 ± 0.06	-
188	14 25 36.30	+53 12 48.03	Q	$0.91084 \pm 0.00020^{\dagger}$	139	1.12	97 ± 20	3.61 ± 0.73	-	II	19.35 ± 0.01	-
189	14 26 03.30	+51 29 37.38	G	$0.41955 \pm 0.15104^*$	138	0.78	471 ± 94	3.01 ± 2.63	0.88 ± 0.09	II	22.88 ± 0.30	-
190	14 27 42.22	+47 18 47.17	G	$0.81534 \pm 0.00045^{\dagger}$	115	0.90	16 ± 3	0.36 ± 0.07	0.41 ± 0.11	II	21.72 ± 0.16	-
191	14 28 57.40	+54 23 52.99	G	$0.40573 \pm 0.00011^{\dagger}$	127	0.71	66 ± 13	0.40 ± 0.08	0.97 ± 0.11	II	18.54 ± 0.02	-
192	14 28 57.66	+54 36 27.81	G	$0.38148 \pm 0.00012^{\dagger}$	200	1.08	251 ± 50	1.46 ± 0.29	1.27 ± 0.11	II	19.16 ± 0.02	-
193	14 29 14.03	+46 56 11.32	G	$0.77764 \pm 0.00027^{\dagger}$	130	0.99	20 ± 4	0.43 ± 0.09	0.45 ± 0.10	III	21.93 ± 0.12	-
194	14 29 33.45	+54 43 35.29	G	$0.12328 \pm 0.00002^{\dagger}$	452	1.03	797 ± 159	0.32 ± 0.06	0.70 ± 0.09	II	15.89 ± 0.00	-
195	14 29 46.05	+51 09 58.42	G	$0.687 \pm 0.161^{\S}$	280	2.04	530 ± 106	12.50 ± 7.58	1.14 ± 0.09	II	21.90 ± 0.14	-
196	14 30 45.29	+45 46 33.57	G	$0.393 \pm 0.042^{\S}$	140	0.77	133 ± 27	0.72 ± 0.23	0.83 ± 0.09	II	19.76 ± 0.02	-
197	14 30 45.65	+50 27 10.84	G	$0.544 \pm 0.058^{\S}$	110	0.72	109 ± 22	1.22 ± 0.40	0.75 ± 0.09	II	20.06 ± 0.05	-
198	14 31 36.99	+52 27 24.90	G	$0.29198 \pm 0.00007^{\dagger}$	186	0.84	371 ± 74	0.97 ± 0.19	0.62 ± 0.09	II	17.57 ± 0.01	-
199	14 31 41.21	+47 44 35.15	Q	$1.17148 \pm 0.00052^{\dagger}$	91	0.77	205 ± 41	13.70 ± 2.78	-	II	20.77 ± 0.04	-
200	14 33 44.95	+48 08 19.45	G	$0.69759 \pm 0.00023^{\dagger}$	108	0.79	350 ± 70	7.21 ± 1.45	0.82 ± 0.09	II	21.63 ± 0.13	-
201	14 34 41.85	+48 03 40.54	Q	$1.61238 \pm 0.00041^{\dagger}$	124	1.08	147 ± 29	20.60 ± 4.21	-	II	20.96 ± 0.04	-
202	14 35 10.3	+49 48 19.49	Q	$0.16606 \pm 0.00005^{\dagger}$	348	1.02	71 ± 14	0.05 ± 0.01	0.44 ± 0.09	II	17.14 ± 0.00	10
203	14 37 16.47	+46 30 28.39	G	$0.683 \pm 0.081^{\S}$	222	1.62	44 ± 9	0.83 ± 0.29	0.73 ± 0.11	III	21.16 ± 0.09	-
204	14 37 30.77	+53 35 42.42	G	$0.50279 \pm 0.00015^{\dagger}$	123	0.78	236 ± 47	2.39 ± 0.48	0.96 ± 0.09	II	20.18 ± 0.04	-
205	14 37 37.66	+51 44 46.45	G	$0.49518 \pm 0.21509^*$	207	1.30	244 ± 49	2.35 ± 2.48	0.93 ± 0.09	II	21.42 ± 0.07	-
206	14 38 04.03	+46 50 51.20	G	$0.74546 \pm 0.00023^{\dagger}$	115	0.87	8 ± 2	0.17 ± 0.04	0.57 ± 0.16	II	21.17 ± 0.10	-
207	14 40 07.85	+55 27 06.08	G	$0.58442 \pm 0.06106^*$	238	1.62	93 ± 19	1.19 ± 0.38	0.69 ± 0.10	II	20.93 ± 0.06	-
208	14 40 13.98	+50 26 01.87	Q	$1.57000 \pm 0.00108^{\dagger}$	143	1.24	136 ± 27	22.00 ± 4.49	0.96 ± 0.10	II	20.65 ± 0.05	-
209	14 40 17.49	+53 26 20.29	G	$0.52672 \pm 0.00013^{\dagger}$	195	1.26	348 ± 70	3.60 ± 0.72	-	II	20.13 ± 0.04	-
210	14 42 15.00	+47 06 46.09	G	$0.44503 \pm 0.00011^{\dagger}$	121	0.71	197 ± 39	1.33 ± 0.27	0.65 ± 0.09	II	19.85 ± 0.03	-
211	14 42 58.59	+52 09 02.07	G	$0.272 \pm 0.057^{\S}$	289	1.24	824 ± 165	1.94 ± 1.02	0.88 ± 0.09	II	17.88 ± 0.01	-
212	14 43 08.22	+51 18 10.01	G	$0.46603 \pm 0.00011^{\dagger}$	152	0.92	7 ± 1	0.06 ± 0.01	<1.10	II	19.63 ± 0.03	-
213	14 43 32.79	+45 46 54.56	G	$0.60401 \pm 0.00023^{\dagger}$	117	0.81	44 ± 9	0.59 ± 0.12	0.61 ± 0.10	II	20.44 ± 0.07	-
214	14 44 07.52	+49 07 30.78	Q	$0.91579 \pm 0.00009^{\dagger}$	193	1.55	175 ± 35	7.59 ± 1.53	0.97 ± 0.09	II	22.01 ± 0.15	-
215	14 44 10.50	+55 47 45.64	G	$0.33475 \pm 0.00005^{\dagger}$	144	0.71	542 ± 108	1.97 ± 0.39	0.72 ± 0.09	II	18.15 ± 0.01	-
216	14 44 58.09	+53 09 29.6	G	$0.38228 \pm 0.13684^*$	167	0.90	184 ± 37	0.82 ± 0.71	0.44 ± 0.09	III	19.47 ± 0.03	-
217	14 45 20.87	+54 03 29.62	G	$0.43345 \pm 0.00013^{\dagger}$	162	0.94	127 ± 26	0.81 ± 0.16	0.66 ± 0.09	II	18.72 ± 0.02	-
218	14 46 07.20	+48 41 37.79	G	$0.37676 \pm 0.00006^{\dagger}$	146	0.78	174 ± 35	0.87 ± 0.17	0.89 ± 0.09	II	19.10 ± 0.02	-
219	14 46 48.98	+47 36 31.43	G	$0.584 \pm 0.052^{\S}$	274	1.86	59 ± 12	0.74 ± 0.22	0.66 ± 0.11	II	20.48 ± 0.05	-
220	14 47 25.82	+56 37 50.13	G	$0.71650 \pm 0.00025^{\dagger}$	105	0.78	12 ± 3	0.23 ± 0.05	<0.50	III	21.32 ± 0.09	-
221	14 50 02.36	+54 05 28.27	G	$0.549 \pm 0.076^{\S}$	126	0.83	507 ± 101	5.78 ± 2.26	-	II	20.51 ± 0.05	-
222	14 50 57.28	+53 00 07.76	Q	$0.91800 \pm 0.00008^{\dagger}$	184	1.48	1407 ± 281	56.50 ± 11.40	0.84 ± 0.09	II	20.59 ± 0.04	-
223	14 51 06.41	+53 33 53.84	Q	$0.43220 \pm 0.00007^{\dagger}$	246	1.43	267 ± 53	1.78 ± 0.36	0.80 ± 0.09	II	18.64 ± 0.01	2
224	14 51 16.59	+51 13 33.28	G	$0.509 \pm 0.050^{\S}$	194	1.23	99 ± 20	0.95 ± 0.29	-	II	19.92 ± 0.05	-
225	14 51 21.28	+56 08 35.45	G	$0.525 \pm 0.073^{\S}$	208	1.34	57 ± 12	0.57 ± 0.22	0.68 ± 0.10	II	20.73 ± 0.06	-
226	14 52 50.73	+51 37 01.90	G	$0.652 \pm 0.126^{\S}$	108	0.77	51 ± 10	0.97 ± 0.50	1.00 ± 0.14	II	21.64 ± 0.11	-
227	14 57 02.81	+48 06 46.64	G	$0.26738 \pm 0.00007^{\dagger}$	168	0.71	5 ± 1	0.01 ± 0.00	<1.11	II	18.00 ± 0.01	-
228	14 57 49.59	+50 45 59.62	G	$0.628 \pm 0.146^{\S}$	199	1.40	55 ± 11	0.90 ± 0.54	0.88 ± 0.11	II	21.58 ± 0.11	-
229	15 00 00.56	+50 43 55.62	G	$0.415 \pm 0.087^{\S}$	142	0.80	12 ± 2	0.07 ± 0.04	<0.83	II	19.70 ± 0.03	-
230	15 01 32.11	+50 34 55.14	G	$0.32062 \pm 0.00007^{\dagger}$	164	0.79	32 ± 6	0.10 ± 0.02	0.73 ± 0.11	II	17.66 ± 0.01	-
231	15 01 48.47	+52 10 47.16	G	$0.492 \pm 0.058^{\S}$	145	0.91	10 ± 2	0.09 ± 0.03	0.67 ± 0.16	II	19.68 ± 0.04	-
232	15 03 33.65	+48 11 17.67	G	$0.672 \pm 0.095^{\S}$	111	0.80	80 ± 16	1.81 ± 0.72	1.16 ± 0.13	II	21.52 ± 0.11	-
233	15 04 45.61	+50 30 08.57	G	$0.65227 \pm 0.05463^*$	201	1.44	1295 ± 259	23.90 ± 6.83	0.91 ± 0.09	II	21.89 ± 0.17	-
234	15 06 12.81	+51 37 07.11	G	$0.61121 \pm 0.00010^{\dagger}$	262	1.82	574 ± 115	8.38 ± 1.69	-	II	20.25 ± 0.05	-
235	15 06 24.10	+53 55 02.61	G	$0.30437 \pm 0.00005^{\dagger}$	184	0.85	118 ± 24	0.33 ± 0.07	0.57 ± 0.09	II	17.95 ± 0.01	-
236	15 18 30.19	+51 58 17.03	G	$0.708 \pm 0.168^{\S}$	105	0.78	51 ± 10	1.17 ± 0.72	0.95 ± 0.11	II	21.84 ± 0.13	-
237	15 18 35.37	+51 04 10.70	G	$0.619 \pm 0.138^{\S}$	119	0.83	371 ± 74	5.63 ± 3.25	0.77 ± 0.09	II	21.14 ± 0.08	-
238	15 18 45.36	+52 37 08.56	G	$0.52105 \pm 0.00011^{\dagger}$	110	0.71	312 ± 62	3.66 ± 0.73	1.11 ± 0.09	II	20.24 ± 0.05	-
239	15 23 11.09	+52 03 03.52	Q	$0.51667 \pm 0.00010^{\dagger}$	132	0.84	382 ± 77	4.35 ± 0.87	1.09 ± 0.09	II	20.74 ± 0.06	-

References. (1) Amirkhanyan et al. (2015); (2) Amirkhanyan (2016); (3) Cotter et al. (1996); (4) Koziel-Wierzbowska & Stasińska (2011); (5) Machalski (1998); (6) Machalski et al. (2001); (7) Nilsson (1998); (8) Rentería Macario & Andernach (2017); (9) Thwala et al. (2019); (10) Kuźmierz et al. (2018)

Table 3: LoTSS GRGs in galaxy clusters: Parameters r_{200} and R_{L*} (cluster richness parameter) have been taken from the WHL galaxy cluster catalog and M_{200} has been computed from Eq.2 of Wen et al. (2012). GRGs marked with † in the redshift column have photometric redshift estimate and the rest have spectroscopic redshifts. r_{200} is the radius within which the mean density of a cluster is 200 times of the critical density of the universe, and M_{200} , is the mass of the cluster within r_{200} . N_{200} is the number of galaxies within r_{200} . Clusters 6, 7 and 20 are identified from GMBCG cluster catalog (Hao et al. 2010).

No	Cluster Name	GRG RA	GRG Dec	z	r_{200}	R_{L*}	M_{200}	N_{200}
					(Mpc)		($10^{14} M_\odot$)	
1	WHLJ112126.4+534457	11 21 26.44	+53 44 56.71	0.10378 ± 0.00002	1.7	69.4	4.6	46
2	WHLJ113250.7+505705	11 32 50.67	+50 57 04.68	0.35857 ± 0.00010	0.9	16.9	0.9	14
3	WHLJ113503.2+482612	11 35 03.20	+48 26 12.12	0.22597 ± 0.00006	1.1	21.7	1.2	17
4	WHLJ121900.8+505255	12 19 00.76	+50 52 54.41	0.38509 ± 0.00008	0.9	17.3	0.9	14
5	GMBCG J188.75636+53.29864	12 35 01.52	+53 17 55.09	0.34480 ± 0.00030	-	-	-	-
6	GMBCG J191.35486+47.11948	12 45 25.16	+47 07 10.11	0.51179 ± 0.00015	-	-	-	-
7	WHLJ124913.2+500044	12 49 13.20	+50 00 43.68	0.34956 ± 0.00010	1.1	25.3	1.4	30
8	WHL J125142.1+503425	12 51 42.04	+50 34 24.65	0.54904 ± 0.00007	1.1	24.3	1.4	11
9	WHL J131404.6+543938	13 14 04.60	+54 39 37.88	0.34645 ± 0.00008	1.2	26.1	1.5	23
10	WHLJ132229.1+504845	13 22 29.07	+50 48 44.53	0.18439 ± 0.00004	0.9	15.9	0.8	11
11	WHLJ133258.3+535356	13 32 58.28	+53 53 55.60	0.32261 ± 0.00009	0.9	16.6	0.9	13
12	WHLJ133512.0+494427	13 35 12.01	+49 44 27.20	$0.436 \pm 0.041^\dagger$	0.9	14.7	0.7	9
13	WHLJ133618.8+533952	13 36 18.75	+53 39 52.12	0.30138 ± 0.00006	0.9	13.4	0.7	11
14	WHLJ134207.0+472553	13 42 06.98	+47 25 53.04	0.17213 ± 0.00004	0.9	13.2	0.7	13
15	WHLJ142857.4+542353	14 28 57.40	+54 23 52.99	0.40573 ± 0.00011	1.2	30.7	1.8	22
16	WHLJ142933.4+544336	14 29 33.45	+54 43 35.29	0.12328 ± 0.00002	1.3	34.7	2.0	29
17	WHLJ144410.5+554746	14 44 10.50	+55 47 45.64	0.33475 ± 0.00005	1.2	30.1	1.7	26
18	WHLJ144520.9+540330	14 45 20.87	+54 03 29.62	0.43345 ± 0.00013	1.0	25.1	1.4	18
19	GMBCG J222.81919+51.22588	14 51 16.59	+51 13 33.28	$0.509 \pm 0.050^\dagger$	-	-	-	-
20	WHLJ150624.1+535503	15 06 24.10	+53 55 02.61	0.30437 ± 0.00005	1.0	19.1	1.0	11

Acknowledgements. We acknowledge the comments and suggestions of the anonymous referee which has helped us improve the quality of our paper. PD and JB gratefully acknowledge generous support from the Indo-French Center for the Promotion of Advanced Research (Centre Franco-Indien pour la Promotion de la Recherche Avancée) under programme no. 5204-2 and thank IUCAA and IUCAA Radio Physics Lab (RPL) for financial and logistic support. This research has made use of the Dutch national e-infrastructure with support of the SURF Cooperative (e-infra 180169) and the LOFAR e-infra group. The Jülich LOFAR Long Term Archive and the German LOFAR network, are both coordinated and operated by the Jülich Supercomputing Centre (JSC), and computing resources on the Supercomputer JUWELS at JSC were provided by the Gauss Centre for Supercomputing e.V. (www.gauss-centre.eu, grant CHTB00) through the John von Neumann Institute for Computing (NIC). HJR gratefully acknowledges generous support from the European Research Council under the European Unions Seventh Framework Programme (FP/2007-2013)/ERC Advanced Grant NEWCLUSTERS-321271. MJH acknowledges support from the UK Science and Technology Facilities Council [ST/R000905/1]. This research has made use of the University of Hertfordshire high-performance computing facility (<https://uhhpc.herts.ac.uk/>) and the LOFAR-UK compute facility, located at the University of Hertfordshire and supported by STFC [ST/P000096/1]. HJR and KJD acknowledges support from the ERC Advanced Investigator programme NewClusters 321271. LOFAR, the Low Frequency Array designed and constructed by ASTRON, has facilities in several countries, which are owned by various parties (each with their own funding), and are collectively operated by the International LOFAR Telescope (ILT) foundation under a joint scientific policy. We thank the LOFAR Galaxy Zoo team. We gratefully acknowledge the use of Edward (Ned) Wright's online Cosmology Calculator. This research has made use of the NASA Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. This research has also made use of the SIMBAD database, operated at CDS, Strasbourg, France. This publication makes use of data products from the Wide-field Infrared Survey Explorer, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration.

References

- Amirkhanyan, V. R., Afanasiev, V. L., & Moiseev, A. V. 2015, *Astrophysical Bulletin*, 70, 45
- Amirkhanyan, V. R. 2016, *Astrophysical Bulletin*, 71, 384
- Bagchi, J., Durret, F., Neto, G. B. L., & Paul, S. 2006, *Science*, 314, 791
- Bagchi, J., Vivek, M., Vikram, V., et al. 2014, *ApJ*, 788, 174
- Becker, R. H., White, R. L., & Helfand, D. J. 1995, *ApJ*, 450, 559
- Begelman, M. C., Blandford, R. D., & Rees, M. J. 1984, *Reviews of Modern Physics*, 56, 255
- Bock, D. C.-J., Large, M. I., & Sadler, E. M. 1999, *AJ*, 117, 1578
- Chambers, K. C., Magnier, E. A., Metcalfe, N., et al. 2016, *arXiv:1612.05560*
- Colless, M., Dalton, G., Maddox, S., et al. 2001, *MNRAS*, 328, 1039
- Condon, J. J. 1992, *ARA&A*, 30, 575
- Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *AJ*, 115, 1693
- Cotter, G., Rawlings, S., & Saunders, R. 1996, *MNRAS*, 281, 1081
- Croston, J. H., Hardcastle, M. J., Mingo, B., et al. 2019, *A&A*, 622, A10.
- Clarke, A. O., Heald, G., Jarrett, T., et al. 2017, *A&A*, 601, A25
- Dabhade, P., Gaikwad, M., Bagchi, J., et al. 2017, *MNRAS*, 469, 2886
- Duncan, K. J., Sabater, J., Röttgering, H. J. A., et al. 2019, *A&A*, 622, A3.
- Fanaroff, B. L., & Riley, J. M. 1974, *MNRAS*, 167, 31P
- Gawroński, M. P., Marecki, A., Kunert-Bajraszewska, M., & Kus, A. J. 2006, *A&A*, 447, 63
- Gopal-Krishna, & Wiita, P. J. 2002, *New A Rev.*, 46, 357
- Gruppioni, C., Zamorani, G., de Ruiter, H. R., et al. 1997, *MNRAS*, 286, 470
- Hardcastle, M. J., Gürkan, G., van Weeren, R. J., et al. 2016, *MNRAS*, 462, 1910
- Hardcastle, M. J., Williams, W. L., Best, P. N., et al. 2019, *A&A*, 622, A12.
- Hao, J., McKay, T. A., Koester, B. P., et al. 2010, *ApJS*, 191, 254
- Hart, R. E., Bamford, S. P., Willett, K. W., et al. 2016, *MNRAS*, 461, 3663.
- Heald, G. H., Pizzo, R. F., Orrú, E., et al. 2015, *A&A*, 582, A123
- Hota, A., Sirothia, S. K., Ohya, Y., et al. 2011, *MNRAS*, 417, L36
- Huchra, J. P., Macri, L. M., Masters, K. L., et al. 2012, *ApJS*, 199, 26
- Hurley-Walker, N., Callingham, J. R., Hancock, P. J., et al. 2017, *MNRAS*, 464, 1146
- Intema, H. T., Jagannathan, P., Mooley, K. P., & Frail, D. A. 2017, *A&A*, 598, A78
- Ishwara-Chandra, C. H., & Saikia, D. J. 1999, *MNRAS*, 309, 100
- Ishwara-Chandra, C. H., Sirothia, S. K., Wadadekar, Y., et al. 2010, *MNRAS*, 405, 436
- Isobe, N., & Koyama, S. 2015, *Publications of the Astronomical Society of Japan*, 67, 77.
- Jamroz, M., Konar, C., Saikia, D. J., et al. 2007, *MNRAS*, 378, 581
- Jennison, R. C., & Das Gupta, M. K. 1953, *Nature*, 172, 996
- Jetha, N. N., Hardcastle, M. J., Ponman, T. J., et al. 2008, *MNRAS*, 391, 1052
- Jones, D. H., Read, M. A., Saunders, W., et al. 2009, *MNRAS*, 399, 683
- Kaiser, N., Aussel, H., Burke, B. E., et al. 2002, *Proc. SPIE*, 4836, 154
- Kaiser, N., Burgett, W., Chambers, K., et al. 2010, *Proc. SPIE*, 7733, 77330E
- Kapahi, V. K., Athreya, R. M., van Breugel, W., et al. 1998, *ApJS*, 118, 275
- Kapińska, A. D., Terentev, I., Wong, O. I., et al. 2017, *AJ*, 154, 253.
- Komberg, B. V., & Pashchenko, I. N. 2009, *Astronomy Reports*, 53, 1086
- Kozieł-Wierzbowska, D. & Stasińska, G. 2011, *MNRAS*, 415, 1013
- Kronberg, P. P., Colgate, S. A., Li, H., & Dufton, Q. W. 2004, *ApJ*, 604, L77
- Kuligowska, E., & Kuźmicz, A. 2018, XXXVIII Polish Astronomical Society Meeting, 82
- Kuźmicz, A., Jamroz, M., Kozieł-Wierzbowska, D., et al. 2017, *Monthly Notices of the Royal Astronomical Society*, 471, 3806
- Kuźmicz, A., Jamroz, M., Bronarska, K., et al. 2018, *The Astrophysical Journal Supplement Series*, 238, 9.
- Lacy, M., Rawlings, S., Saunders, R., et al. 1993, *MNRAS*, 264, 721
- Laing, R. A., Riley, J. M., & Longair, M. S. 1983, *MNRAS*, 204, 151
- Lara, L., Cotton, W. D., Feretti, L., et al. 2001, *A&A*, 370, 409
- Lara, L., Giovannini, G., Cotton, W. D., et al. 2004, *A&A*, 421, 899
- Ledlow, M. J., & Owen, F. N. 1996, *AJ*, 112, 9
- Lopes, P. A. A. 2007, *MNRAS*, 380, 1608.
- Lynden-Bell, D. 1969, *Nature*, 223, 690
- Machalski, J. 1998, *A&AS*, 128, 153
- Machalski, J., Jamroz, M., & Zola, S. 2001, *A&A*, 371, 445
- Machalski, J., Kozieł-Wierzbowska, D., & Jamroz, M. 2007, *Acta Astron.*, 57, 227
- Mack, K.-H., Klein, U., O'Dea, C. P., Willis, A. G., & Saripalli, L. 1998, *A&A*, 329, 431
- Mahatma, V. H., Hardcastle, M. J., Williams, W. L., et al. 2019, *A&A*, 622, A13.
- Malarecki, J. M., Jones, D. H., Saripalli, L., Staveley-Smith, L., & Subrahmanyam, R. 2015, *MNRAS*, 449, 955
- McGilchrist, M. M., & Riley, J. M. 1990, *MNRAS*, 246, 123
- Miley, G., & De Breuck, C. 2008, *A&A Rev.*, 15, 67
- Mohan, N., & Rafferty, D. 2015, *Astrophysics Source Code Library*, ascl:1502.007
- Mahony, E. K., Morganti, R., Prandoni, I., et al. 2016, *MNRAS*, 463, 2997
- Nandi, S., & Saikia, D. J. 2012, *Bulletin of the Astronomical Society of India*, 40, 121
- Nandi, S., Saikia, D. J., Roy, R., et al. 2019, *MNRAS*, 486, 5158
- Nilsson, K. 1998, *A&AS*, 132, 31
- Oort, M. J. A., Steemers, W. J. G., & Windhorst, R. A. 1988, *A&AS*, 73, 103
- O'Sullivan, S. P., Machalski, J., Van Eck, C. L., et al. 2019, *A&A*, 622, A16.
- Pâris, I., Petitjean, P., Aubourg, É., et al. 2018, *A&A*, 613, A51.
- Paul, S., John, R. S., Gupta, P., & Kumar, H. 2017, *MNRAS*, 471, 2
- Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, *A&A*, 594, A13
- Prescott, M., Whittam, I. H., Jarvis, M. J., et al. 2018, *MNRAS*, 480, 707
- Richards, G. T., Myers, A. D., Gray, A. G., et al. 2009, *The Astrophysical Journal Supplement Series*, 180, 67

- Richards, G. T., Myers, A. D., Peters, C. M., et al. 2015, *The Astrophysical Journal Supplement Series*, 219, 39
- Rengelink, R. B., Tang, Y., de Bruyn, A. G., et al. 1997, *A&AS*, 124, 259
- Rentería Macario, J., & Andernach, H. 2017, arXiv:1710.10731
- Röttgering, H. J. A., Wieringa, M. H., Hunstead, R. W., & Ekers, R. D. 1997, *MNRAS*, 290, 577
- Saikia, D. J., Thomasson, P., Jackson, N., Salter, C. J., & Junor, W. 1996, *MNRAS*, 282, 837
- Saikia, D. J., & Jamrozy, M. 2009, *Bulletin of the Astronomical Society of India*, 37, 63
- Safouris, V., Subrahmanyan, R., Bicknell, G. V., & Saripalli, L. 2009, *MNRAS*, 393, 2
- Saripalli, L., Hunstead, R. W., Subrahmanyan, R., & Boyce, E. 2005, *AJ*, 130, 896
- Saripalli, L., & Malarecki, J. M. 2015, *The Many Facets of Extragalactic Radio Surveys: Towards New Scientific Challenges*, 44
- Schoenmakers, A. P., de Bruyn, A. G., Röttgering, H. J. A., van der Laan, H., & Kaiser, C. R. 2000, *MNRAS*, 315, 371
- Schoenmakers, A. P., de Bruyn, A. G., Röttgering, H. J. A., & van der Laan, H. 2001, *A&A*, 374, 861
- Sebastian, B., Ishwara-Chandra, C. H., Joshi, R., et al. 2018, *MNRAS*, 473, 4926
- Shimwell, T. W., Röttgering, H. J. A., Best, P. N., et al. 2017, *A&A*, 598, A104
- Shimwell, T. W., Tasse, C., Hardcastle, M. J., et al. 2019, *A&A*, 622, A1.
- Smith, D. J. B., Best, P. N., Duncan, K. J., et al. 2016, *SF2A-2016: Proceedings of the Annual meeting of the French Society of Astronomy and Astrophysics*, 271
- Solovyov, D. I., & Verkhodanov, O. V. 2014, *Astrophysical Bulletin*, 69, 141
- Subrahmanyan, R., Saripalli, L., & Hunstead, R. W. 1996, *MNRAS*, 279, 257
- Thwala, S. A., Shafi, N., Colafrancesco, S., et al. 2019, *MNRAS*, 485, 1938
- van Weeren, R. J., Brüggen, M., Röttgering, H. J. A., et al. 2011, *A&A*, 533, A35
- Wen, Z. L., Han, J. L., & Liu, F. S. 2012, *ApJS*, 199, 34
- Wiita, P. J., Rosen, A., Gopal-Krishna, & Saripalli, L. 1989, *Hot Spots in Extragalactic Radio Sources*, 327, 173
- Williams, W. L., Hardcastle, M. J., Best, P. N., et al. 2019, *A&A*, 622, A2.
- Willis, A. G., Strom, R. G., & Wilson, A. S. 1974, *Nature*, 250, 625
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, 140, 1868
- York, D. G., Adelman, J., Anderson, J. E., Jr., et al. 2000, *AJ*, 120, 1579

Appendix A: Appendix

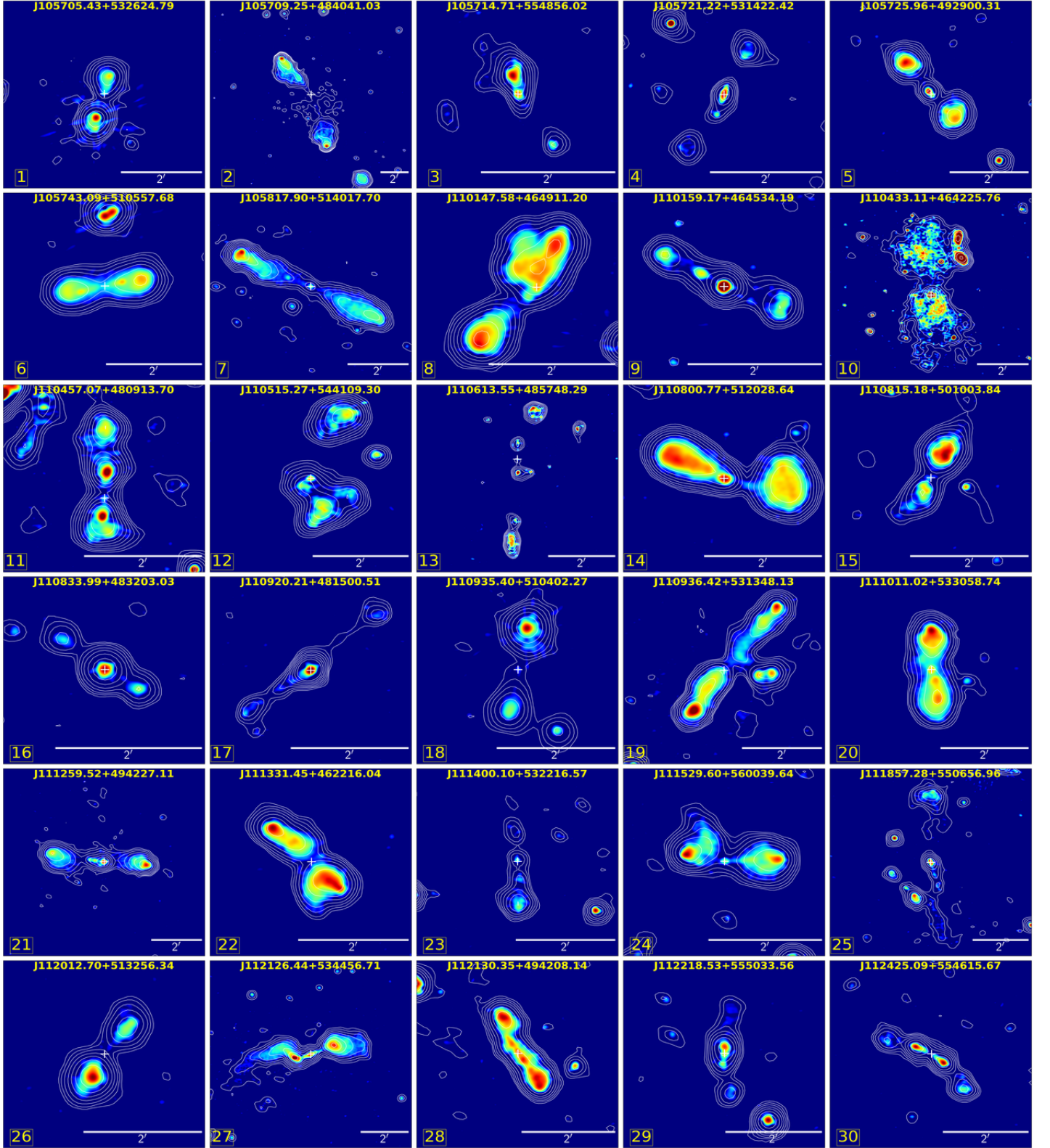


Fig. A.1: LoTSS high resolution ($\sim 6''$) images 1-30 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.

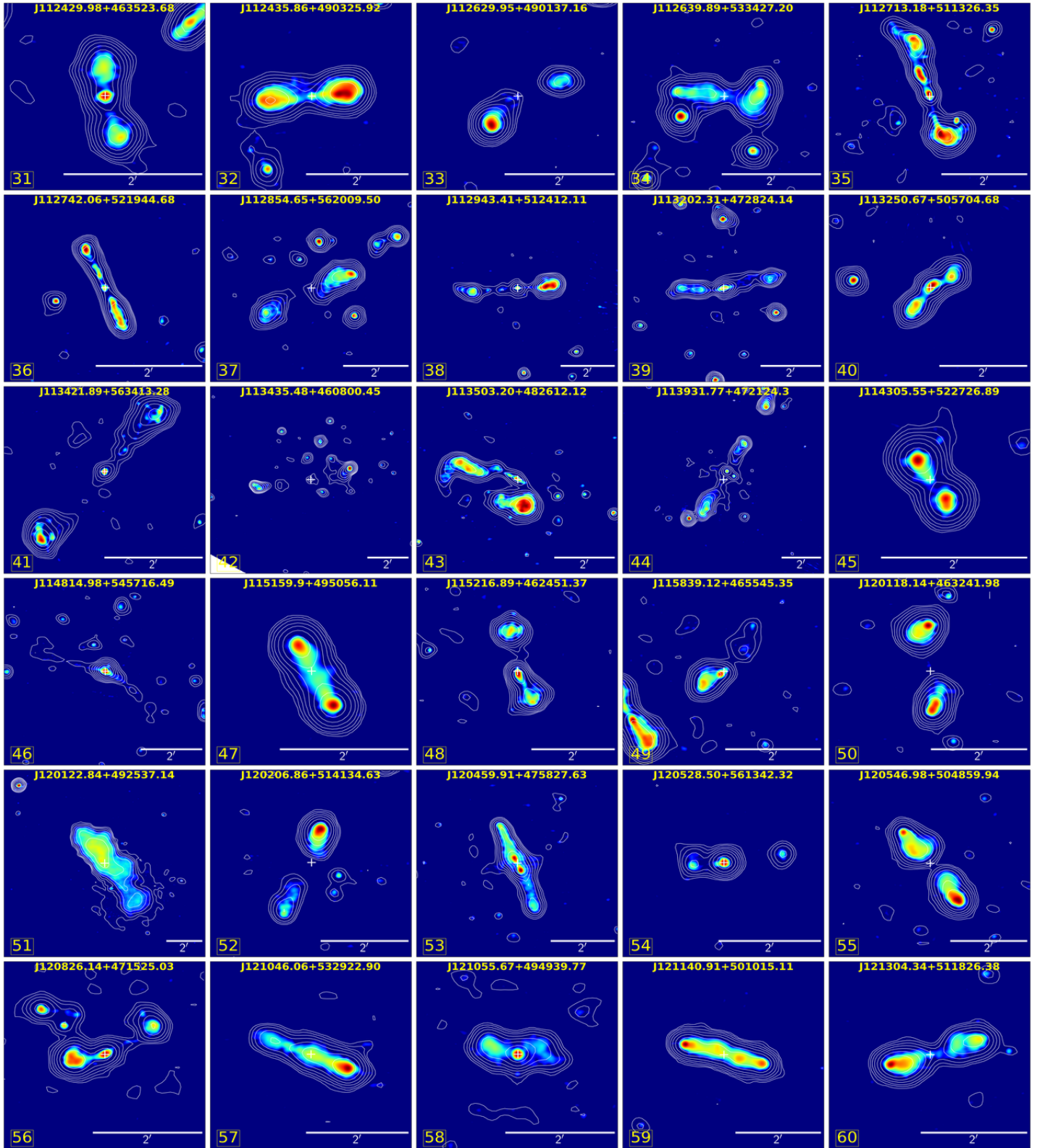


Fig. A.2: LoTSS high resolution ($\sim 6''$) images 31-60 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.

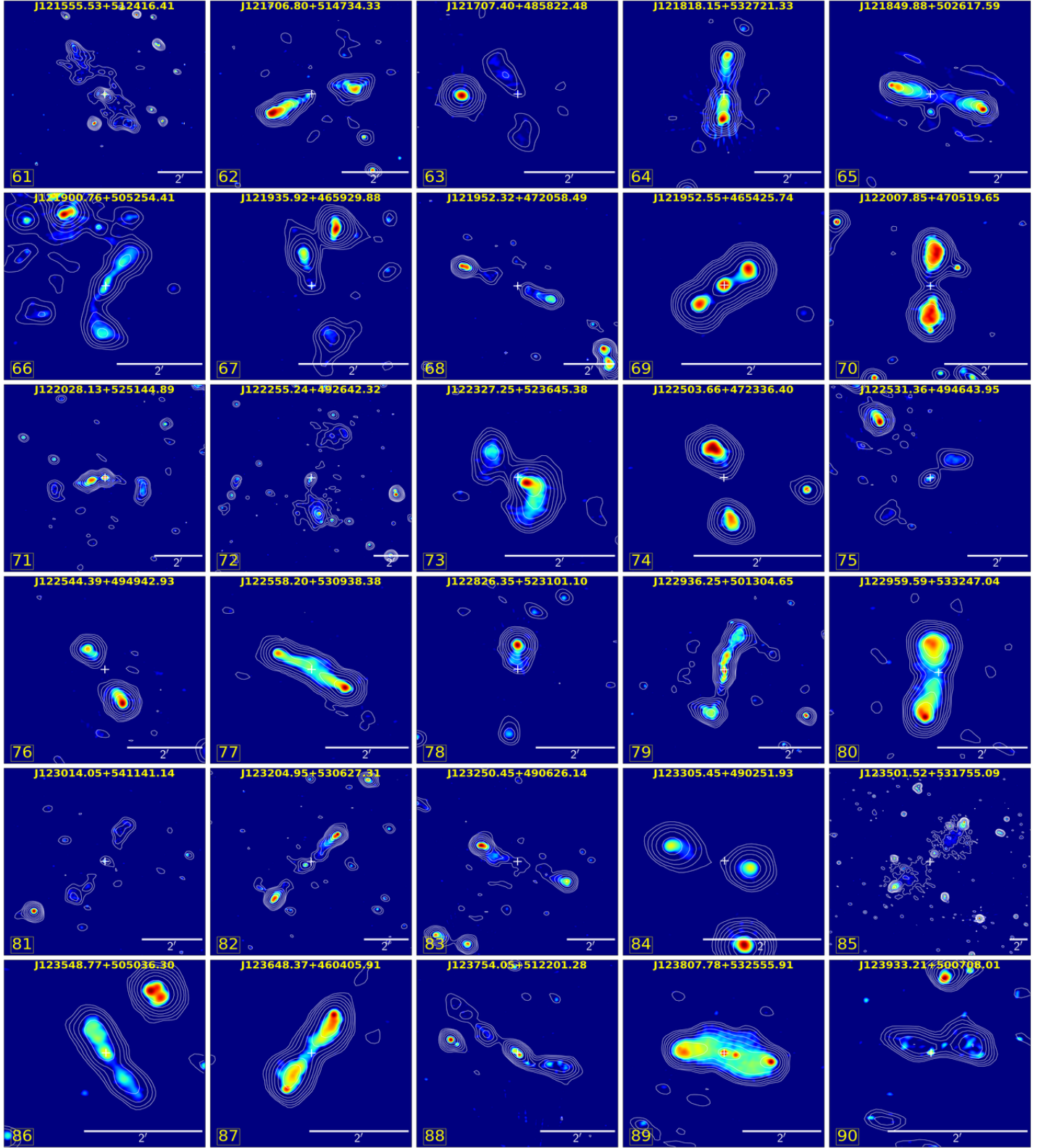


Fig. A.3: LoTSS high resolution ($\sim 6''$) images 61-90 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.

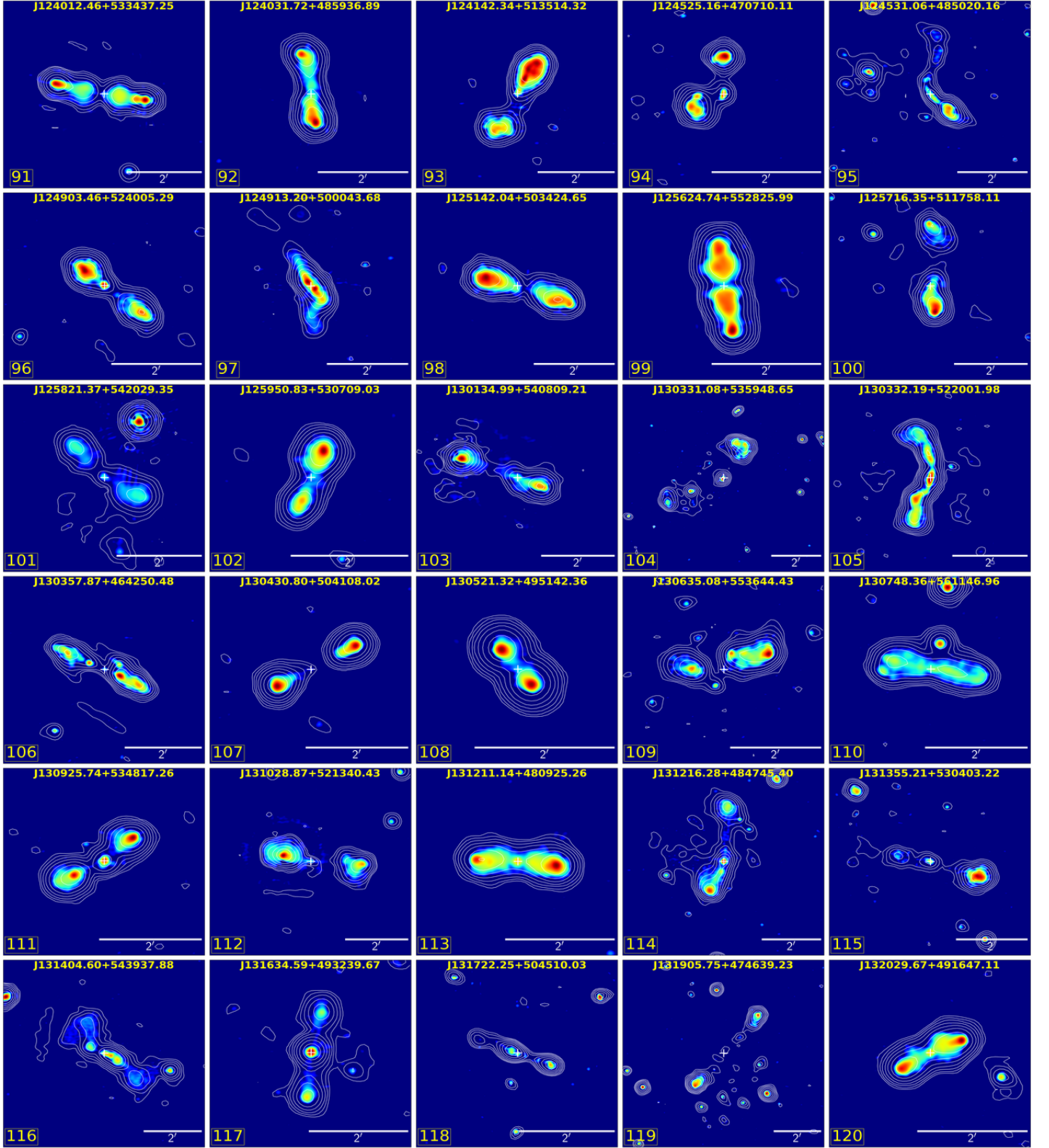


Fig. A.4: LoTSS high resolution ($\sim 6''$) images 91-120 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.

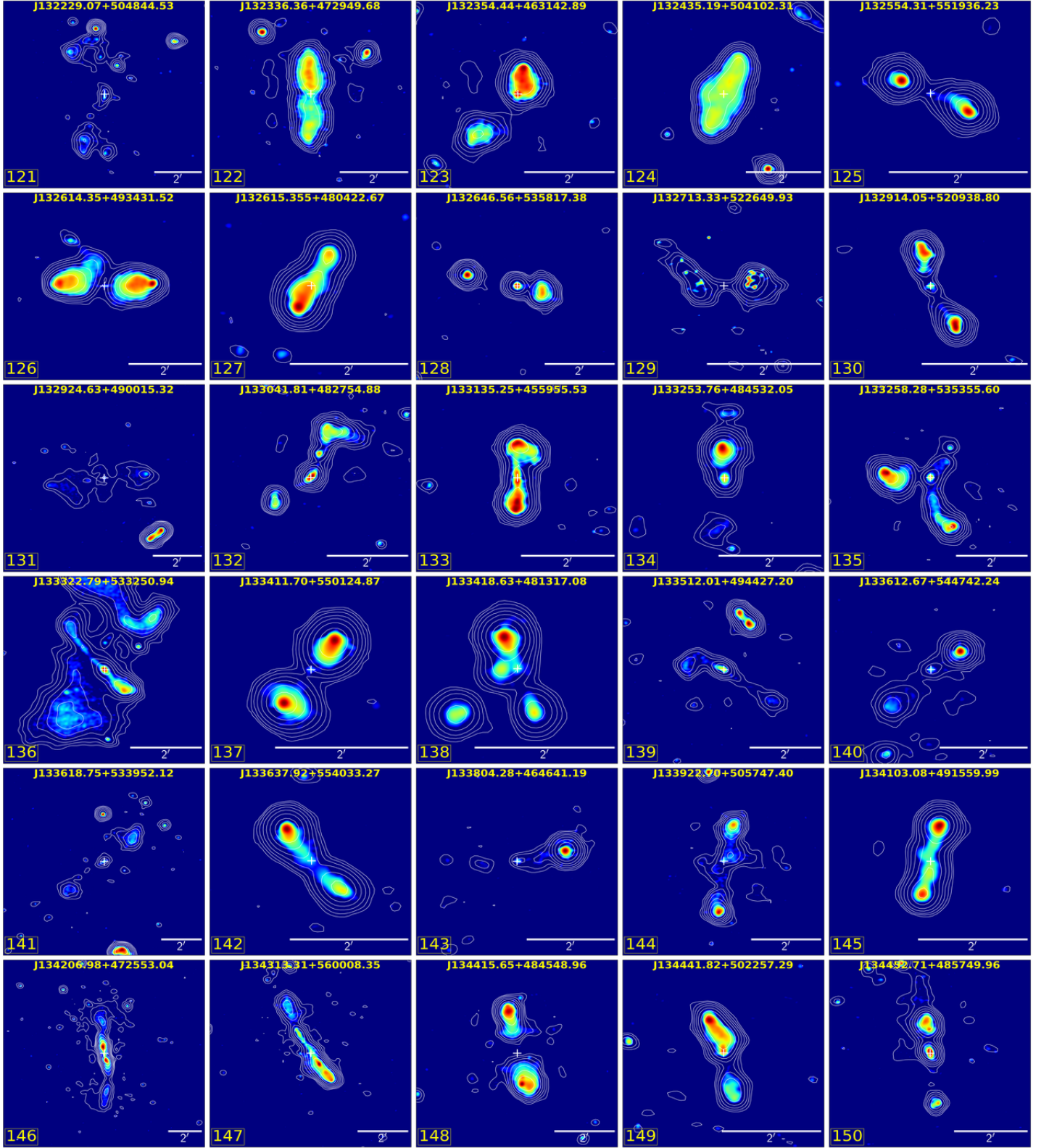


Fig. A.5: LoTSS high resolution ($\sim 6''$) images 121-150 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.

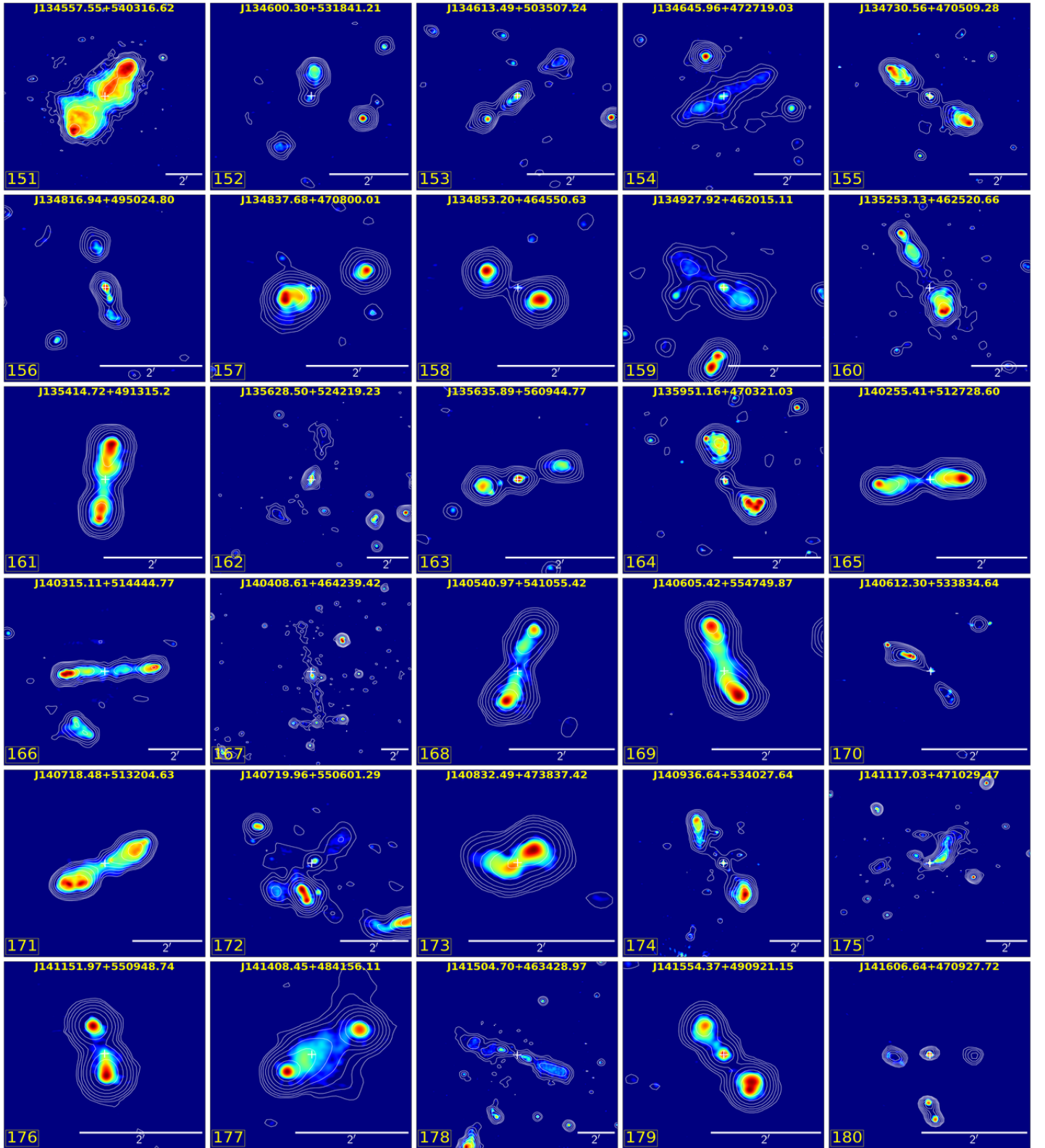


Fig. A.6: LoTSS high resolution ($\sim 6''$) images 151-180 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.

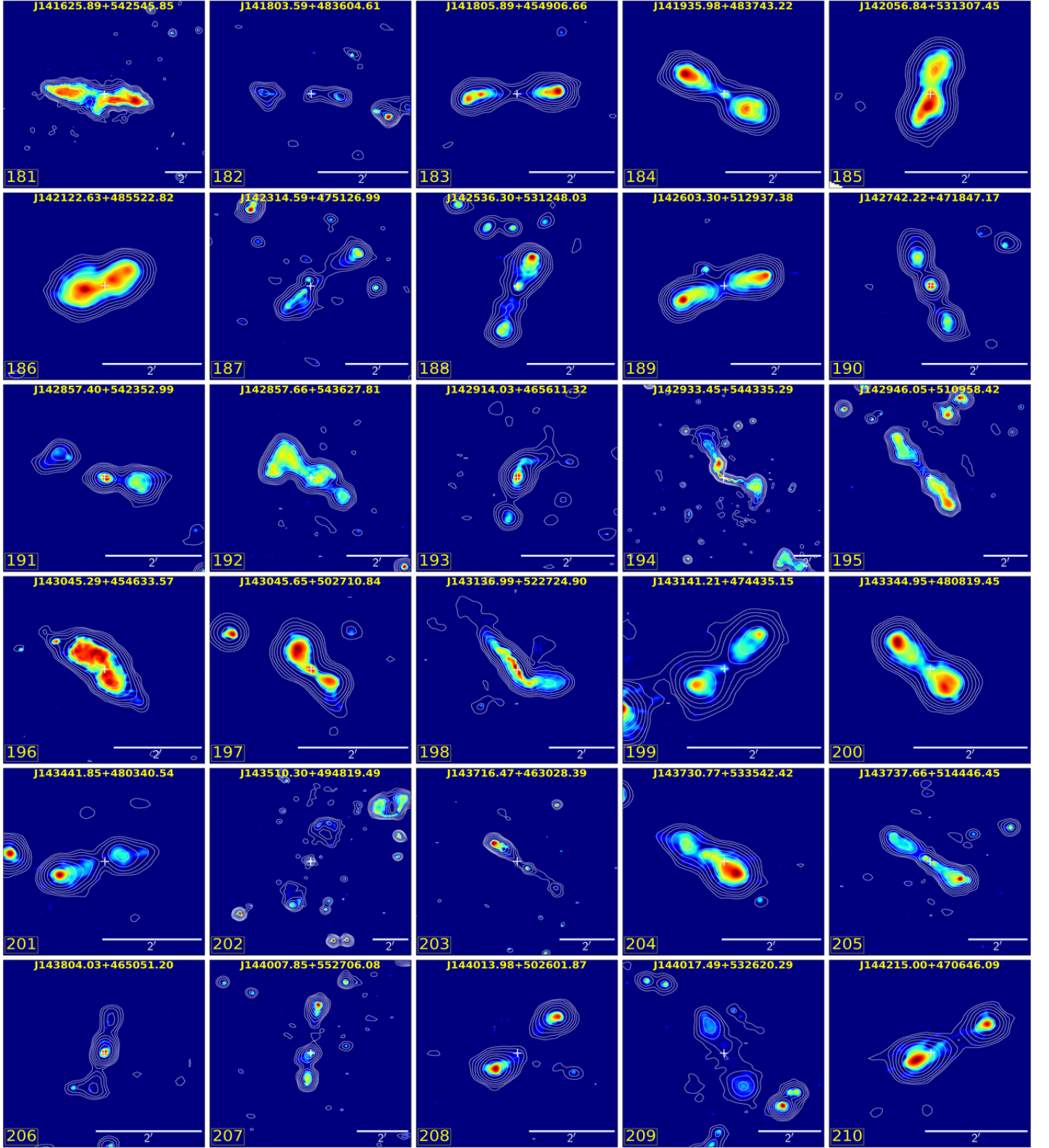


Fig. A.7: LoTSS high resolution ($\sim 6''$) images 181-210 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.

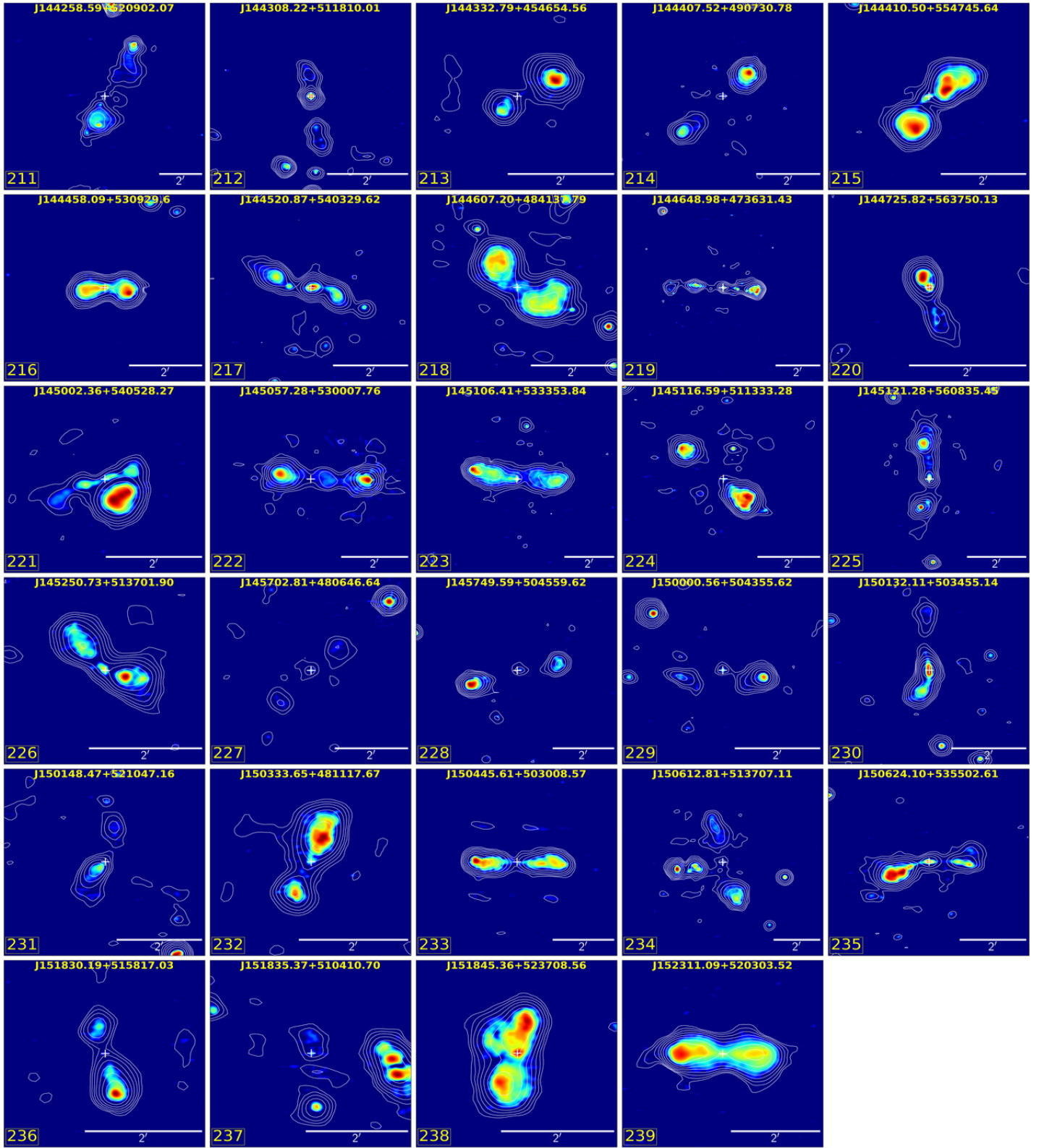


Fig. A.8: LoTSS high resolution ($\sim 6''$) images 211-239 with LoTSS low resolution ($\sim 20''$) contours overlaid. The cross '+' marker indicates the location of the host galaxy. The radio contours are drawn with 8 levels which are chosen by equally (log scale) dividing the data value range above $\sim 3\sigma$, where σ is the local RMS of the map.